Multi-Messenger Astronomy with Cen A

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astro-ph/0805.2608

Outline of the talk

1. Auger correlation claim
2. Test by multi-messenger approach?
   - Cen A – source & acceleration models
   - Our simulation
   - Results
3. Summary and outlook

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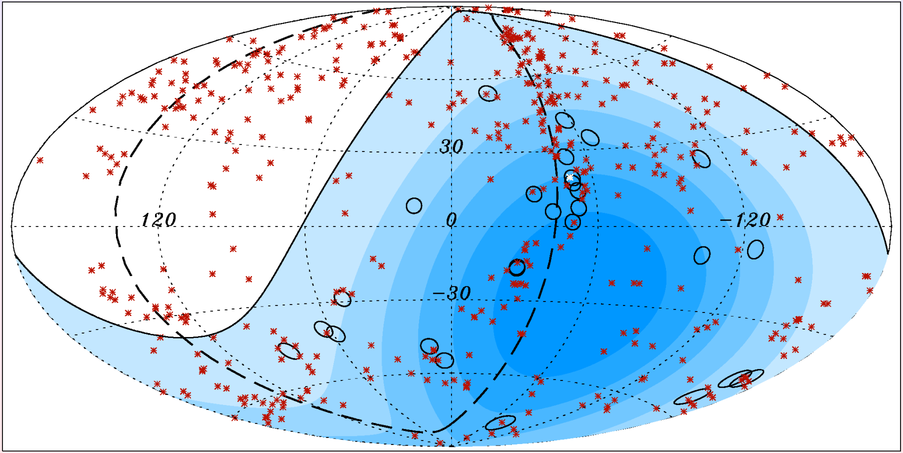
Correlations with AGNs: PAO analysis

- 27 CRs (⊙) and 472 AGN (∗):



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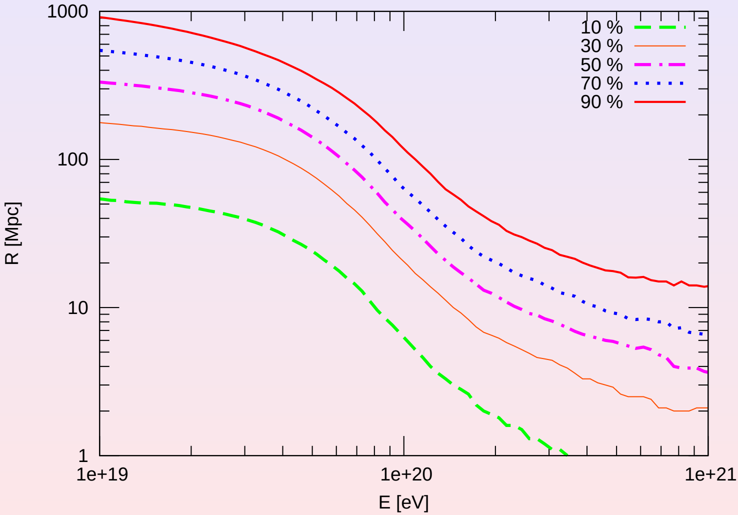
- large fraction of all AGN required to accelerate to $E > 10^{20}$eV
- internal inconsistencies: energy and chemical composition

Energy threshold consistent with GZK horizon?

- 8 out of 13 CRs ($E \geq 57\,\text{EeV}$) correlated within 75 Mpc:



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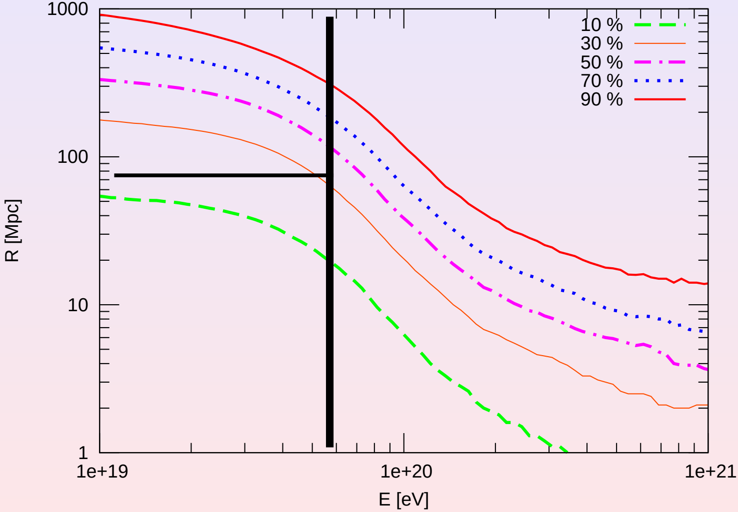

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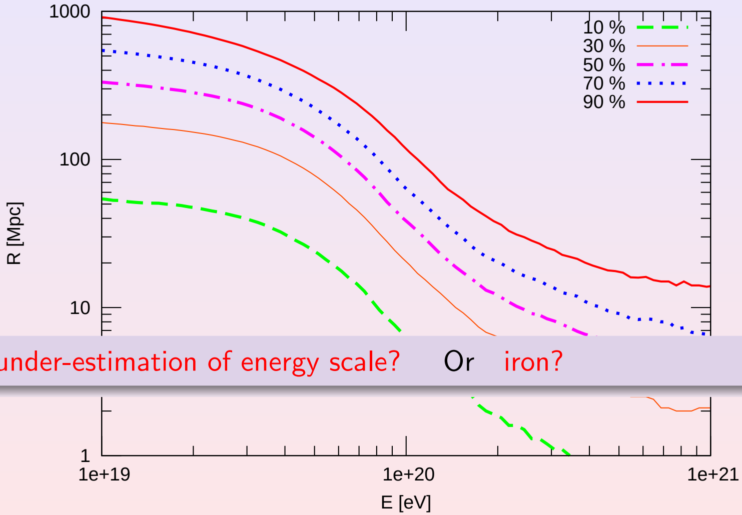

under-estimation of energy scale? Or iron?

Correlations with AGNs: Problems

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- angular scale ℓ consistent with expected deflections?
- confusion danger with other sources in supergalactic plane ($\ell \approx 3^\circ$)
- large fraction of all AGN required to accelerate to $E > 10^{20}$eV
- internal inconsistencies:
  - energy scale
  - chemical composition
- independent/additional evidence?

Possible source/acceleration scenarios

- mechanism: shock acceleration vs. acceleration in regular fields
- location: core, hot spots, along the jet
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- neglect acceleration
- fix 2 basic scenarios: "core" and "jet"

Fixing the source parameters:

- observations:
  - $d = 3.8\,\mathrm{kpc}$
  - $M = (0.5-2)\times 10^8 M_\odot$
  - $\dot{M} = 6\times 10^{-4} M_\odot$
  - $L_X = 5\times 10^{41}\mathrm{erg/s}$

$\Rightarrow$ efficiency $\eta = 5\%$

- supports standard thin, optical thick accretion disc with

$$T(r) = \left(\frac{3GM\dot{M}}{8\sigma\pi r^3}\left[1-(R_0/r)^{1/2}\right]\right)^{1/4}$$

- add X-ray from hot corona
- simplify to 1-dim geometry

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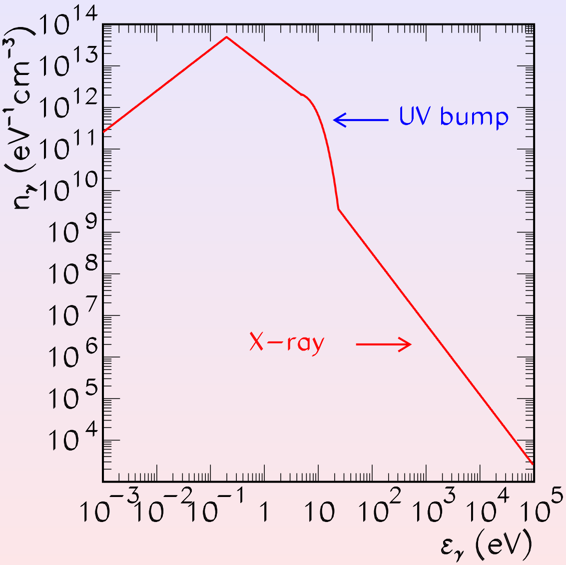
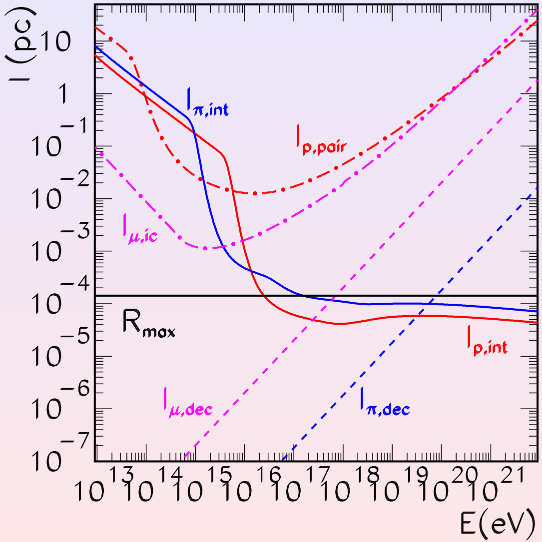
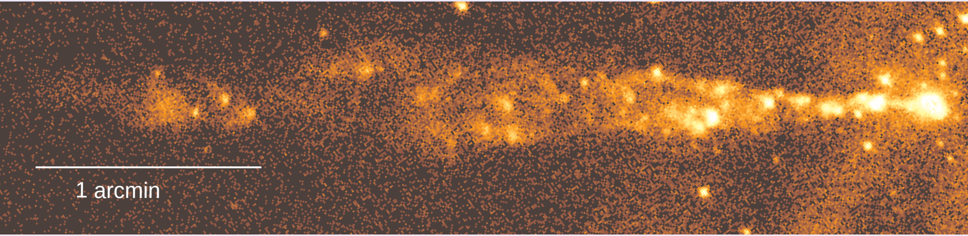
UV and X-ray background from the accretion disk



Lenght scales for acceleration close to the core

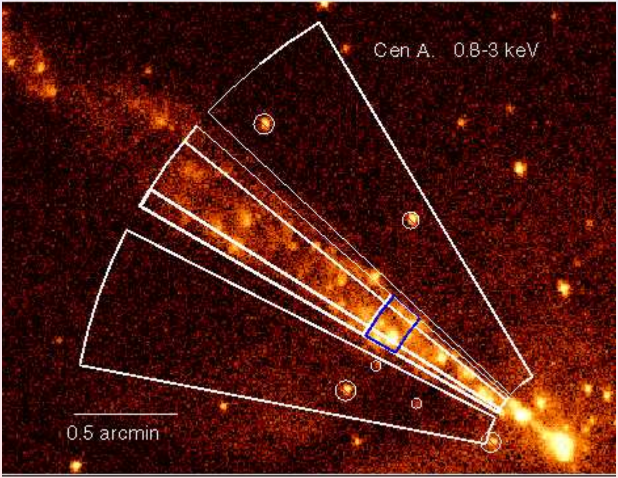


Chandra observation of X-ray emission in the jet

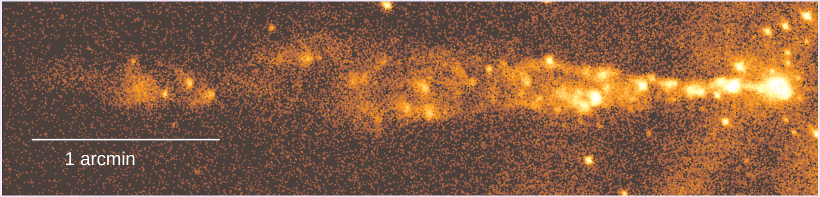


Chandra observation of X-ray emission in the jet

- divide in subareas
- separate fit to gas column density X and spectral index α

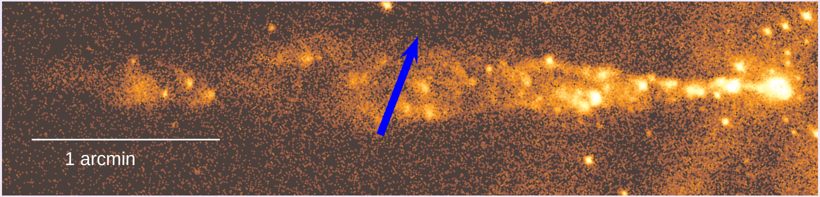


Chandra observation of X-ray emission in the jet: Results



- $X = 1.5 \times 10^{21}/\text{cm}^2$ in the jet

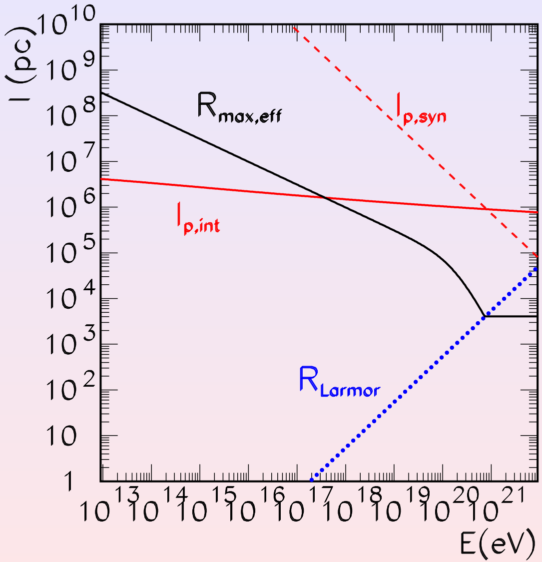
Chandra observation of X-ray emission in the jet: Results



- $X = 1.5 \times 10^{21}/\text{cm}^2$ in the jet
- with $d = 0.4\,\text{kpc}$ and $\sigma_{pp} = 150\text{mbarn}$:

$\Rightarrow$ interaction depth $\tau_{pp} \sim 0.01$

Length scales for acceleration in the jet



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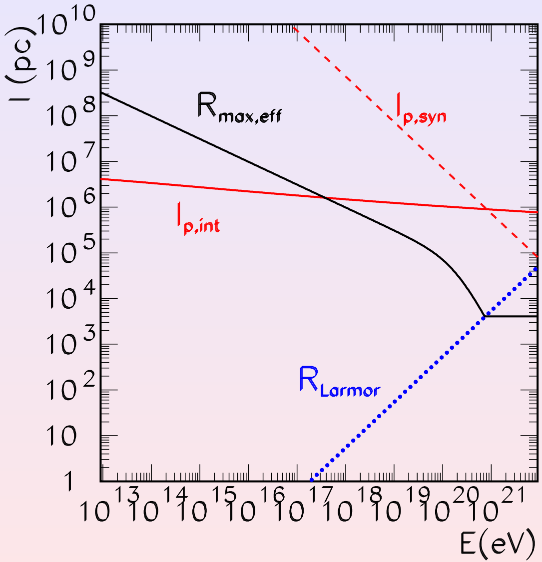

- diffusion increases effective size

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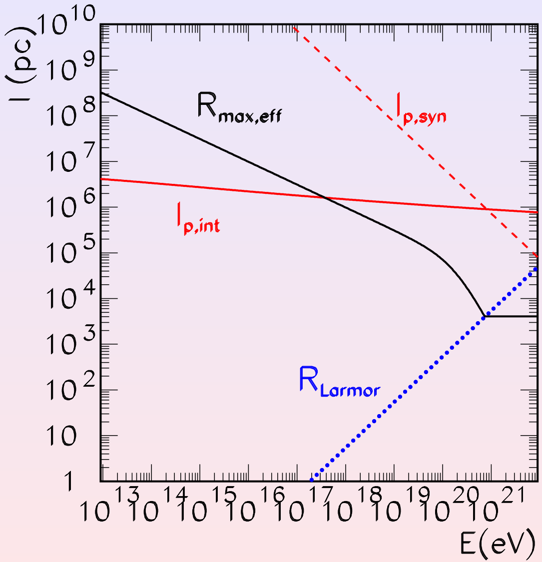


- diffusion increases effective size
- for pp no threshold
- $\tau = 1$ for $E = 10^{17}$eV, optimal for neutrino telescope

Our two base models

acceleration close to the core

acceleration in accretion shock/regular fields

$p\gamma$ interactions

$\tau_{\gamma\gamma} \gg 1$, synchrotron losses for $e^{\pm}$

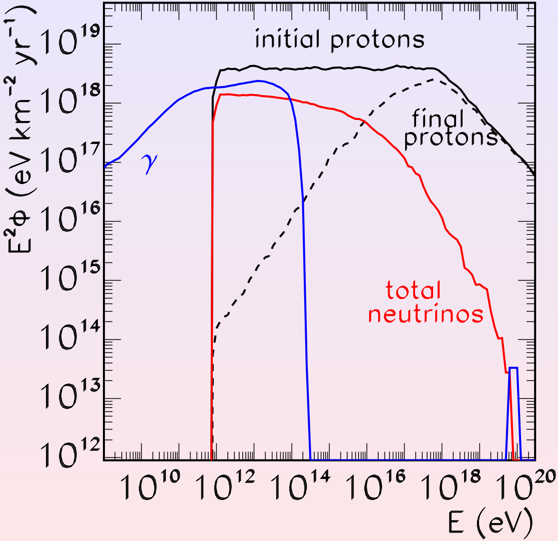
acceleration in jet

shock acceleration

pp interactions

$\tau_{\gamma\gamma} \ll 1$, synchrotron losses for $e^{\pm}$

Results for acceleration in jet: broken power-law



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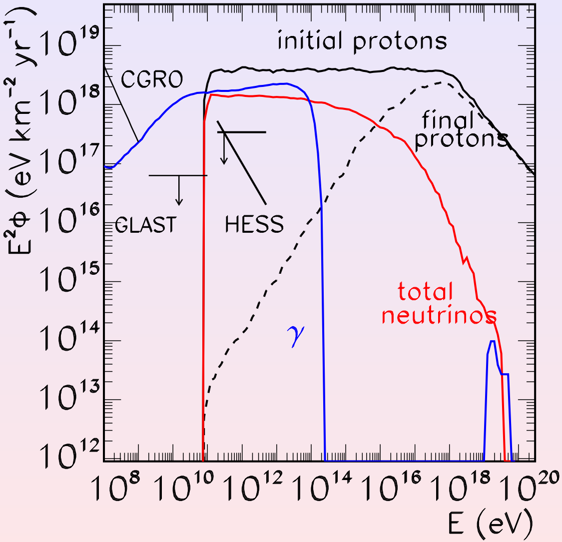

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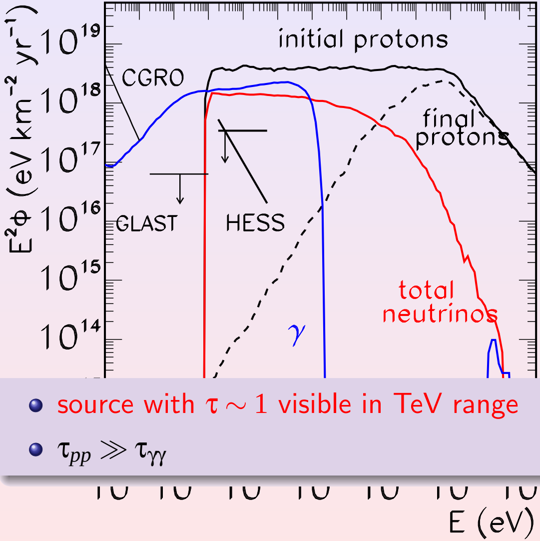

- source with $\tau \sim 1$ visible in TeV range
- $\tau_{pp} \gg \tau_{\gamma\gamma}$

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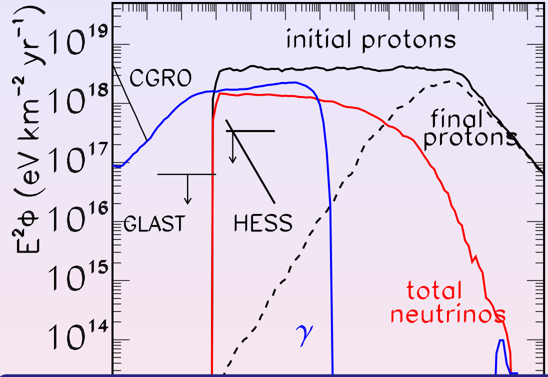


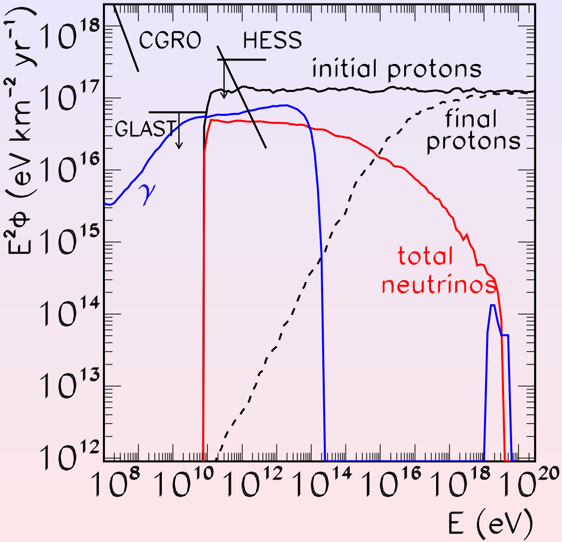
$\alpha = 2.7$ required for diffuse CR flux in "dip model"

- disfavoured as spectrum of single source Cen A

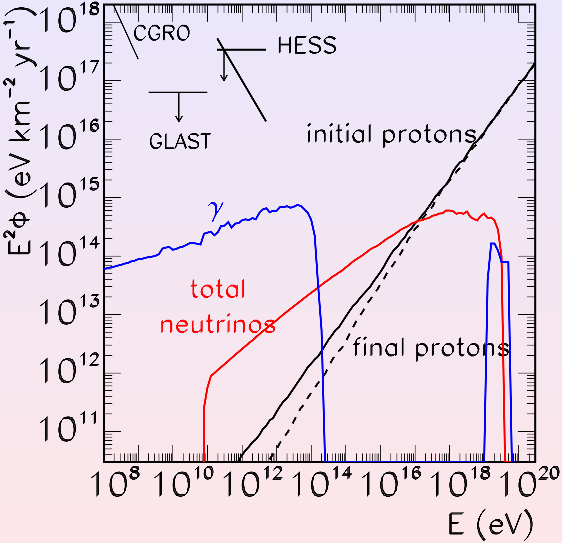
$\Rightarrow$ diffuse spectrum = superposition of single sources with $dn/dE_{\max}$ distribution

- HE γ observations constrain UHECR models

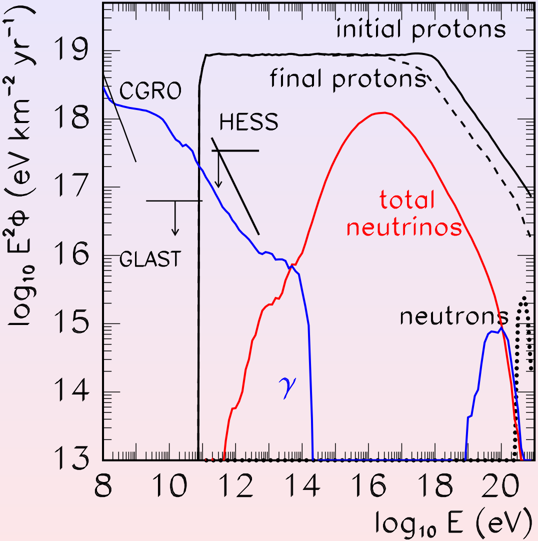
Results for acceleration in jet: $\alpha = 2$



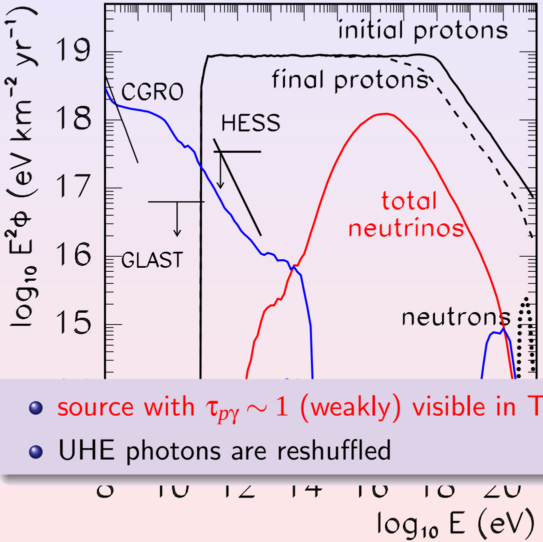
Results for acceleration in jet: $\alpha = 1.2$



Results for acceleration close to the core: broken power-law



Results for acceleration close to the core: broken power-law



- source with $\tau_{p\gamma} \sim 1$ (weakly) visible in TeV range
- UHE photons are reshuffled

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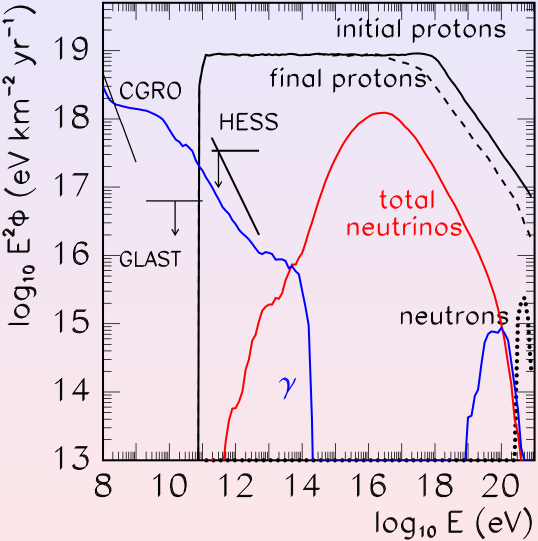

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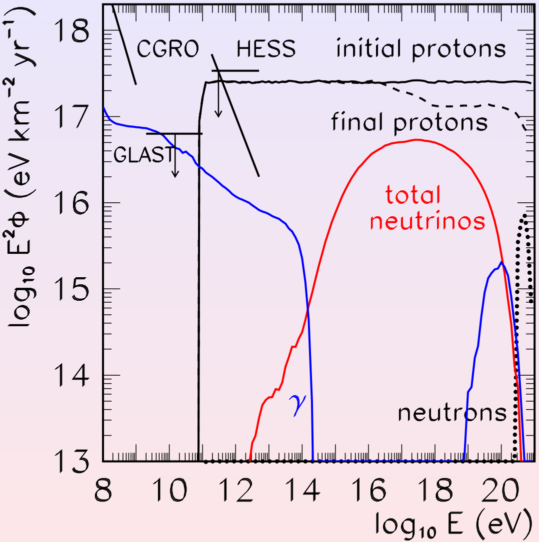
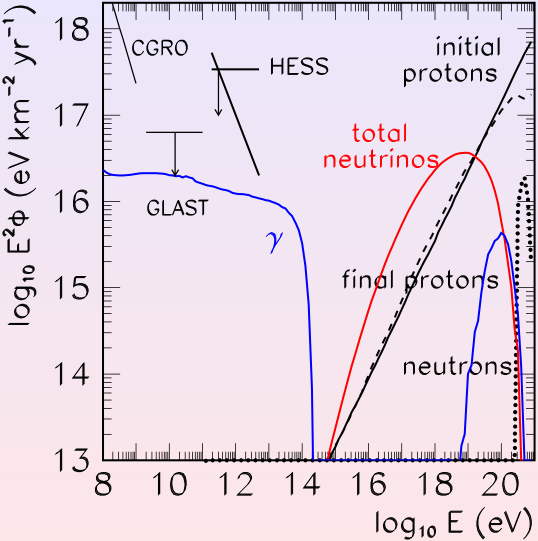
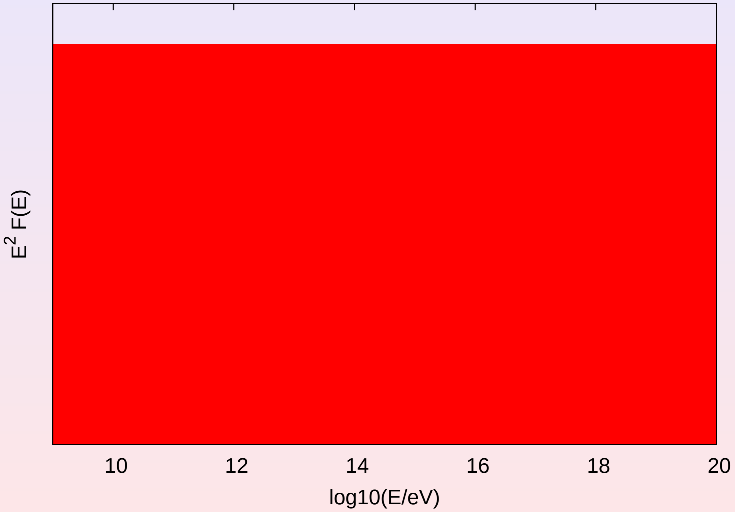


Results for acceleration close to the core: $\alpha = 1.2$



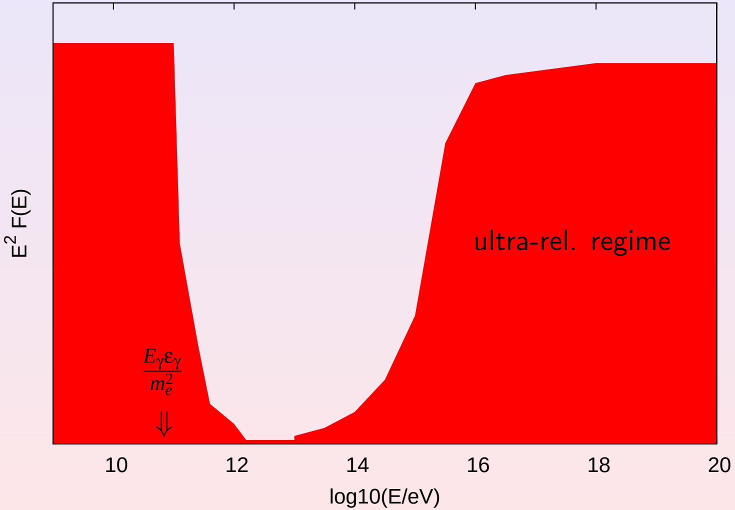
Regenerating TeV photons: a) in the source

- injections spectrum $F_\gamma(E) \propto 1/E^2$



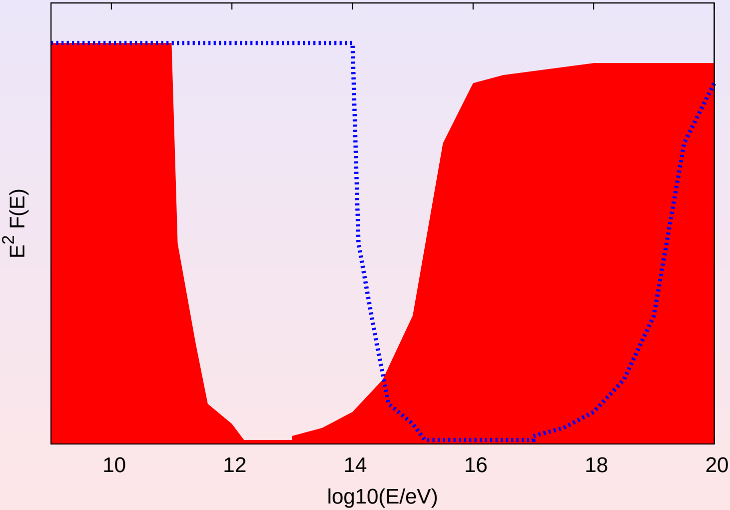
Regenerating TeV photons: a) in the source

- : thin above 10^{16}eV, ultra-rel. regime



Regenerating TeV photons: b) on CMB

- photons above 10^{16}eV cascade on CMB



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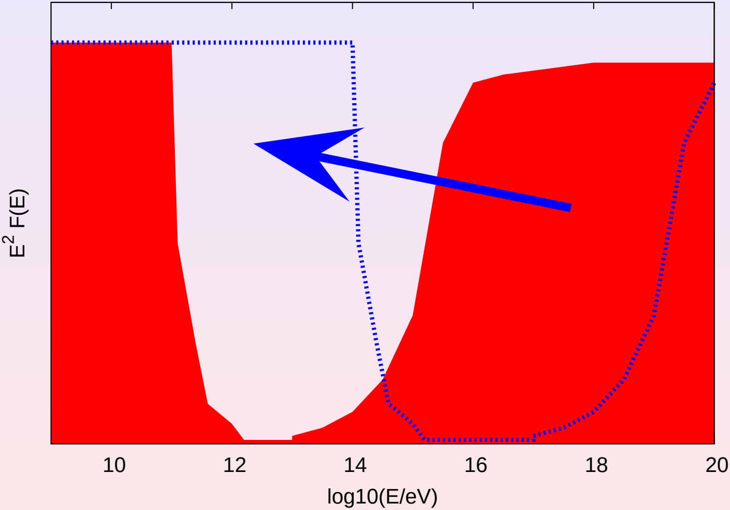

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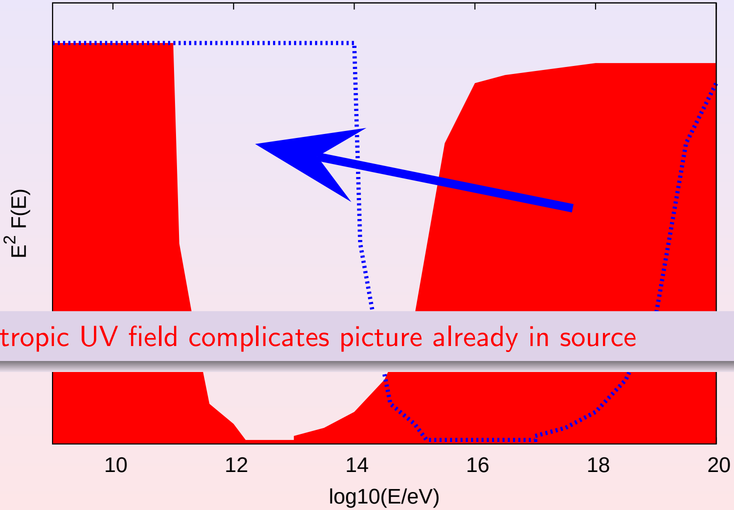

anisotropic UV field complicates picture already in source

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  - pp may be more important than pγ in jet