

A person wearing a blue jacket is holding a silver, reflective horn. The background features a sunset over a body of water, with a complex, golden, geometric pattern resembling a Voronoi diagram or a similar mathematical structure overlaid on the sky. The overall scene is dimly lit, with the primary light source being the setting sun.

# Listening for the Echoes of Inflation with BICEP2 and Beyond

**Jeffrey P. Filippini**

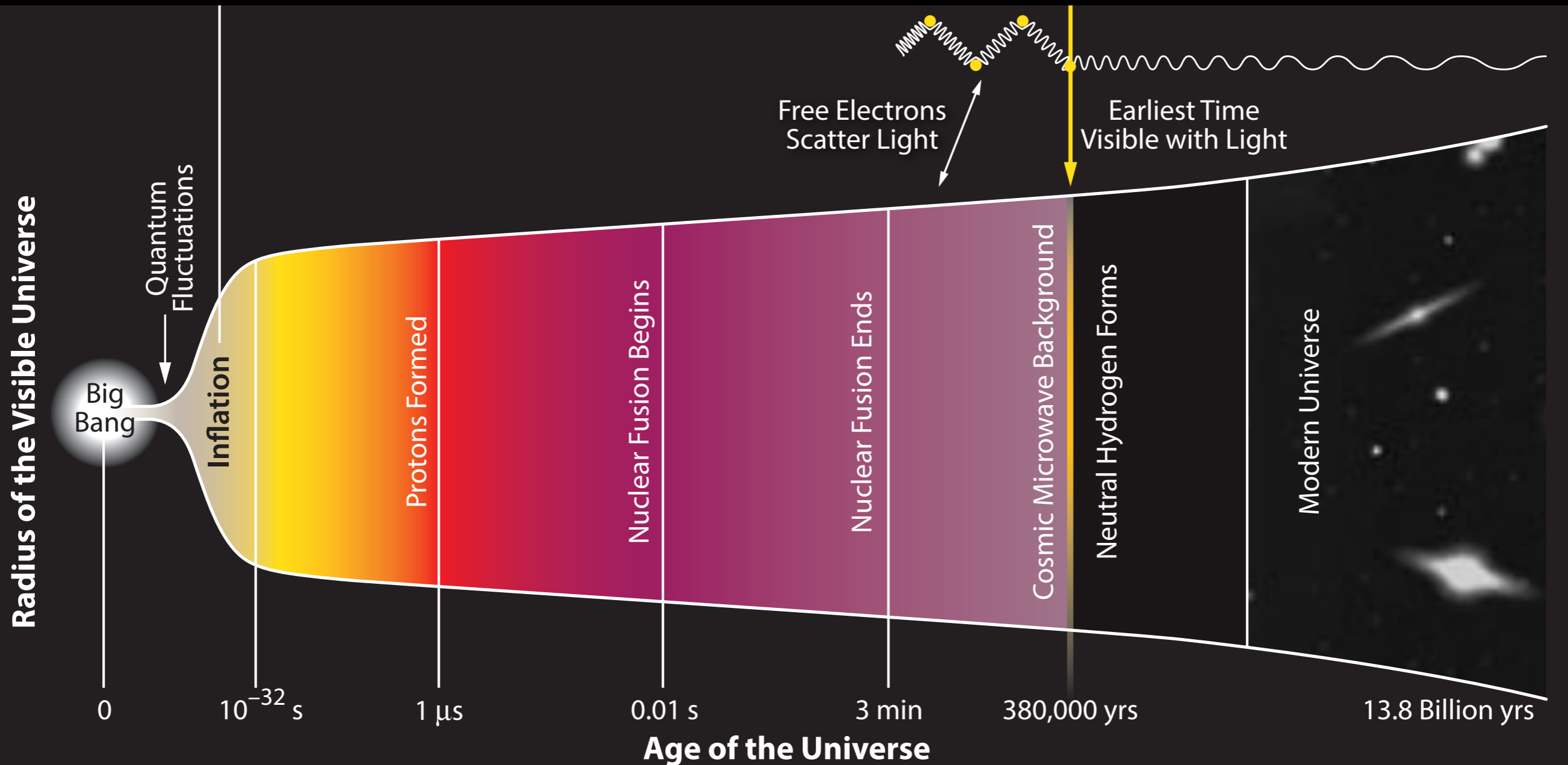
California Institute of Technology  
*for the BICEP2 collaboration*

# Outline

- The Echoes of Inflation
- The BICEP2 Instrument
- Results and Interpretation
- The Future

# A History of Creation

Universe becomes transparent

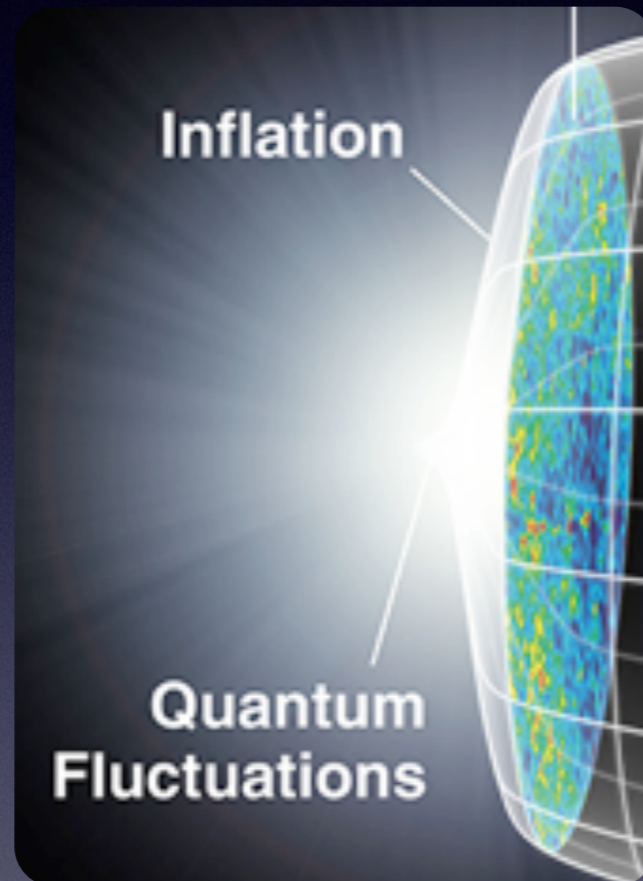


# Inflation

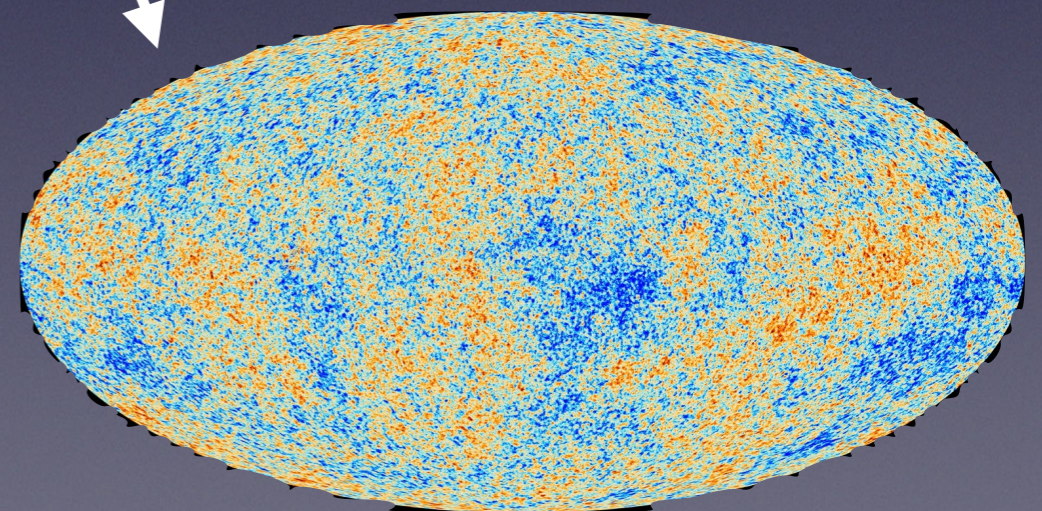
Quantum  
fluctuations...

“**inflaton**”

... imprinted onto  
cosmic scales

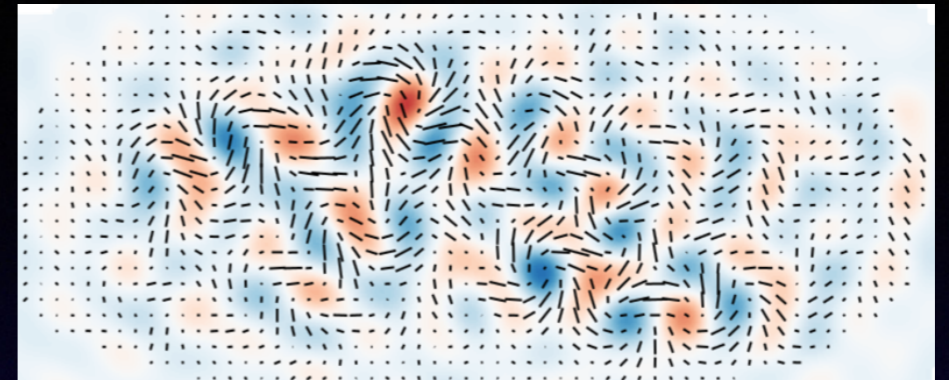


- ✓ Homogeneity
- ✓ Isotropy
- ✓ Nearly-flat geometry ( $\Omega \sim 1$ )
- ✓ Super-horizon fluctuations
- ✓ Nearly-scale invariant density perturbations ( $n_s \sim 1$ )



Density perturbations

# Primordial Perturbations



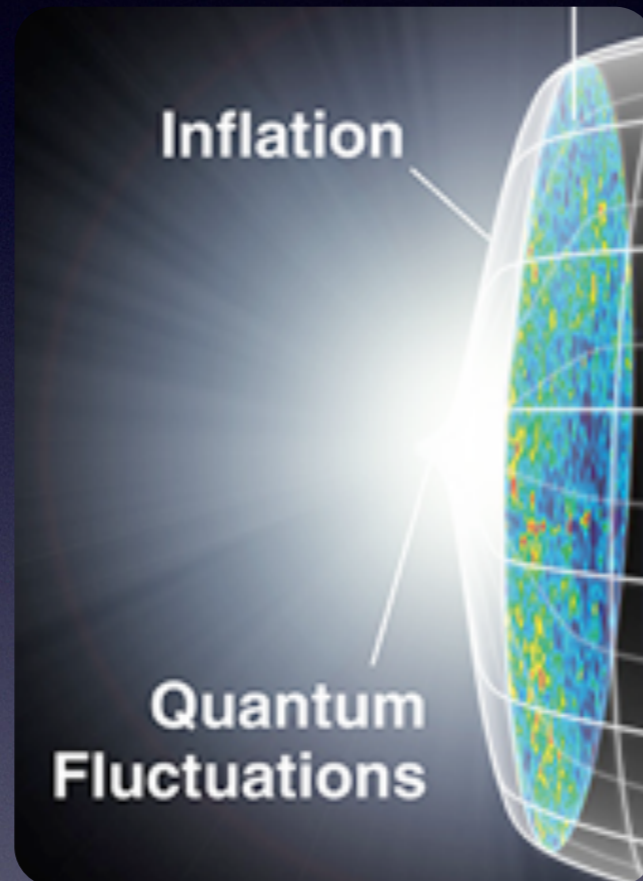
**Primordial gravitational waves**

Quantum fluctuations...

“**inflaton**”

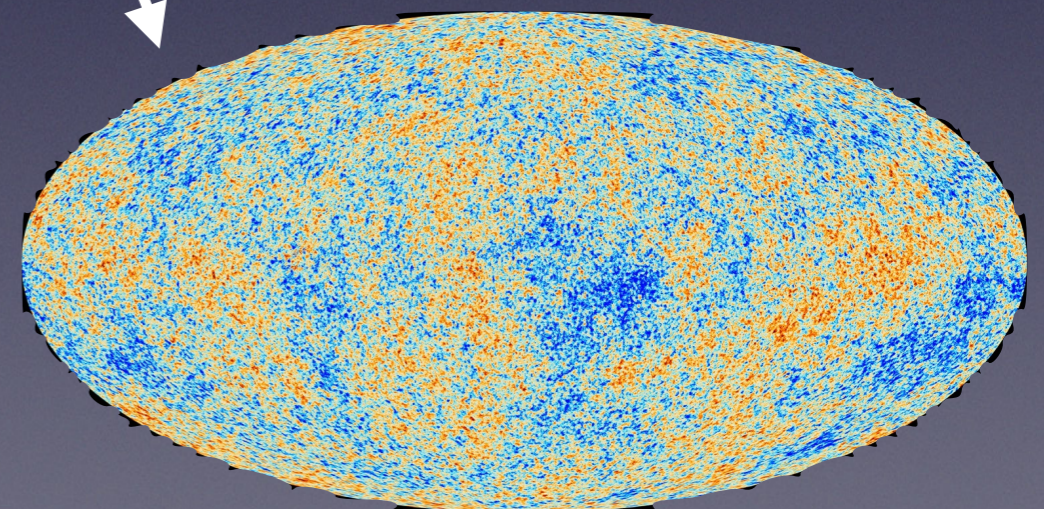
**metric tensor**

... imprinted onto  
cosmic scales



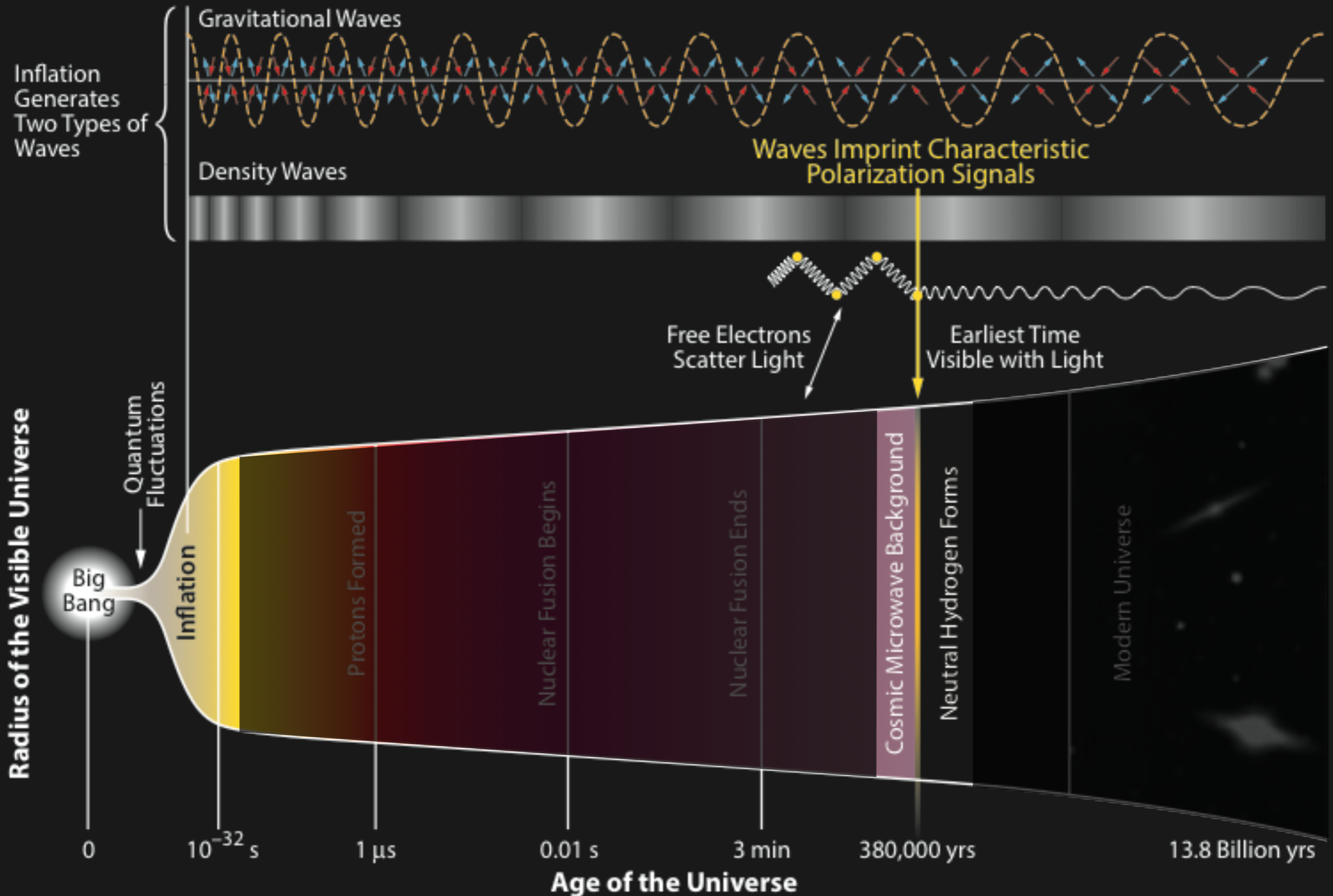
$$r \approx \frac{V[\phi]}{(4 \times 10^{16} \text{ GeV})^4}$$

*GUT-scale physics!?*

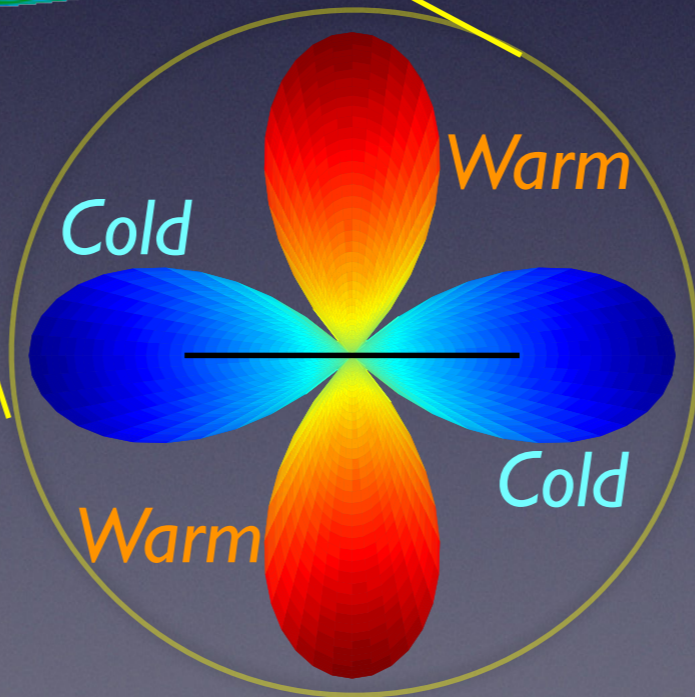
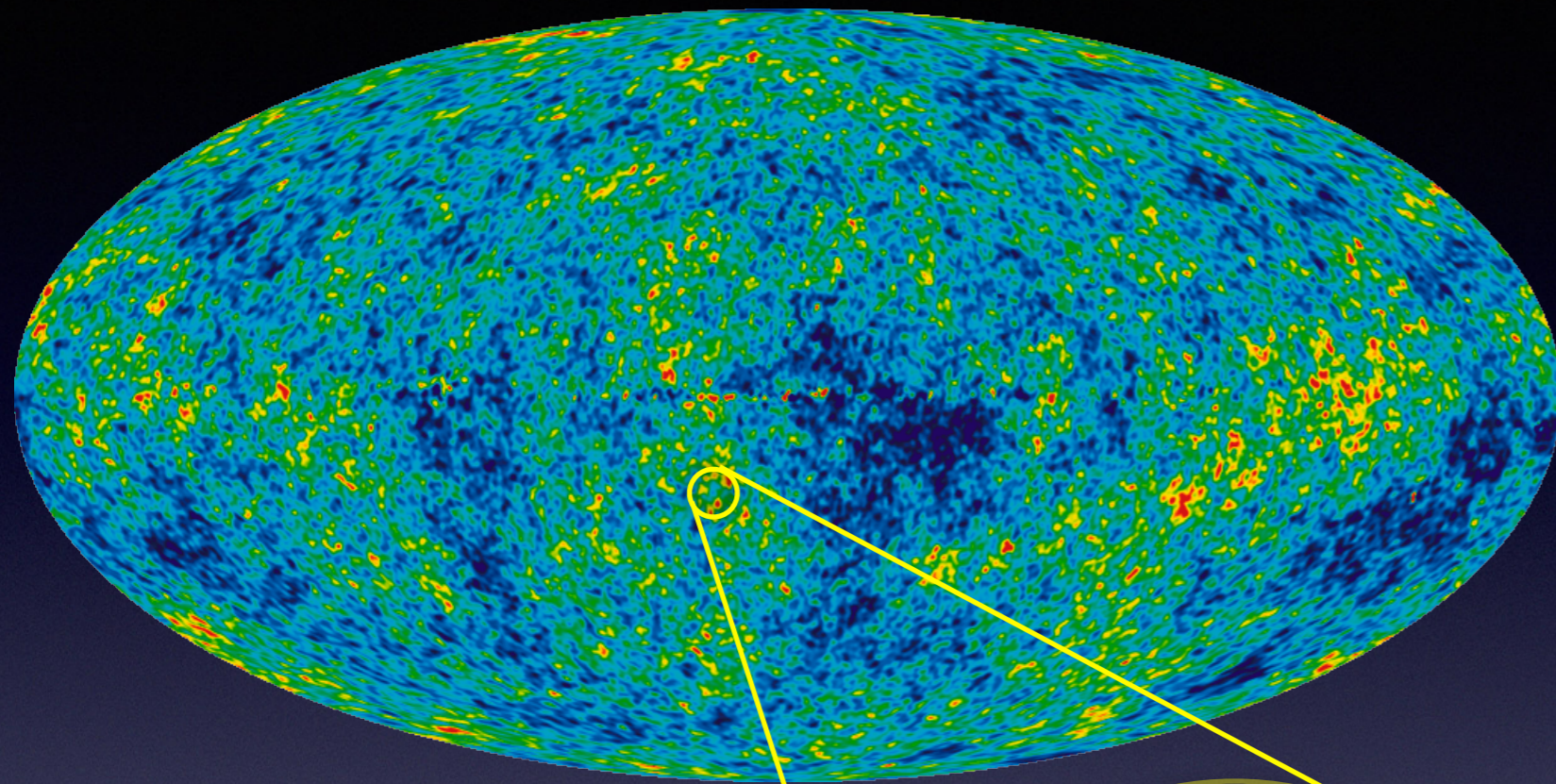


**Density perturbations**

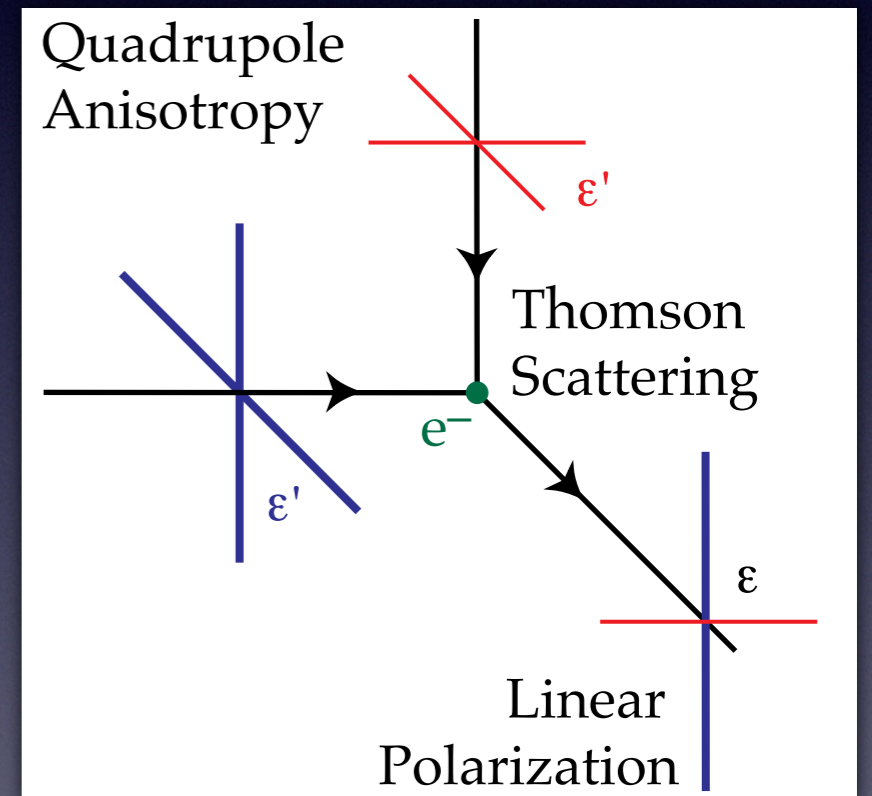
# History of the Universe



# CMB Polarization



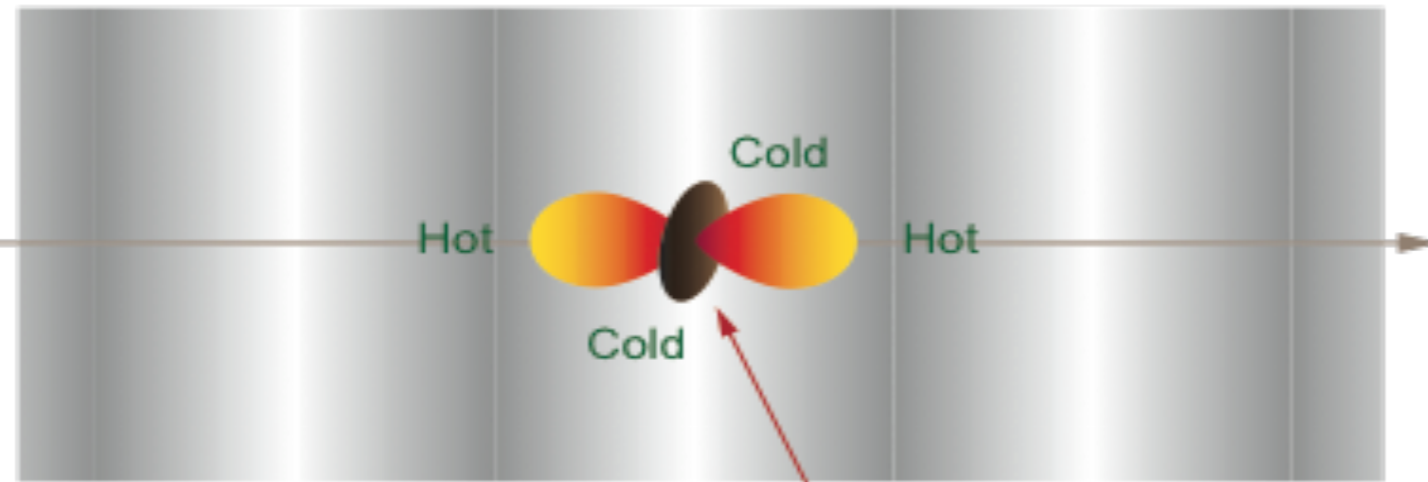
From Thomson scattering wherever there is ionized gas and quadrupole anisotropy



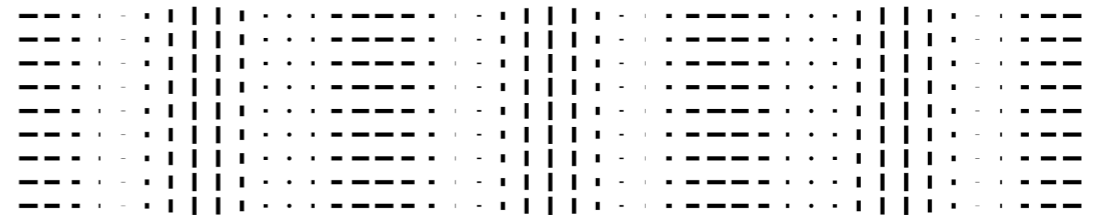
Wayne Hu

# Patterns of Polarization

Density Wave

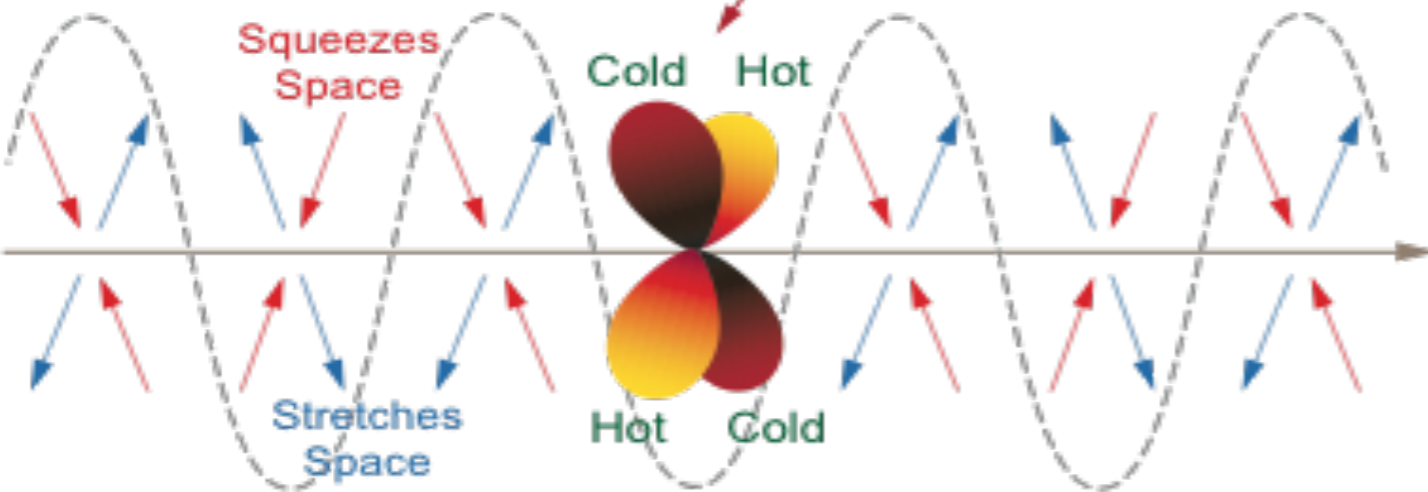


## E-mode Polarization



Temperature Pattern Seen by Electrons

Gravitational Wave



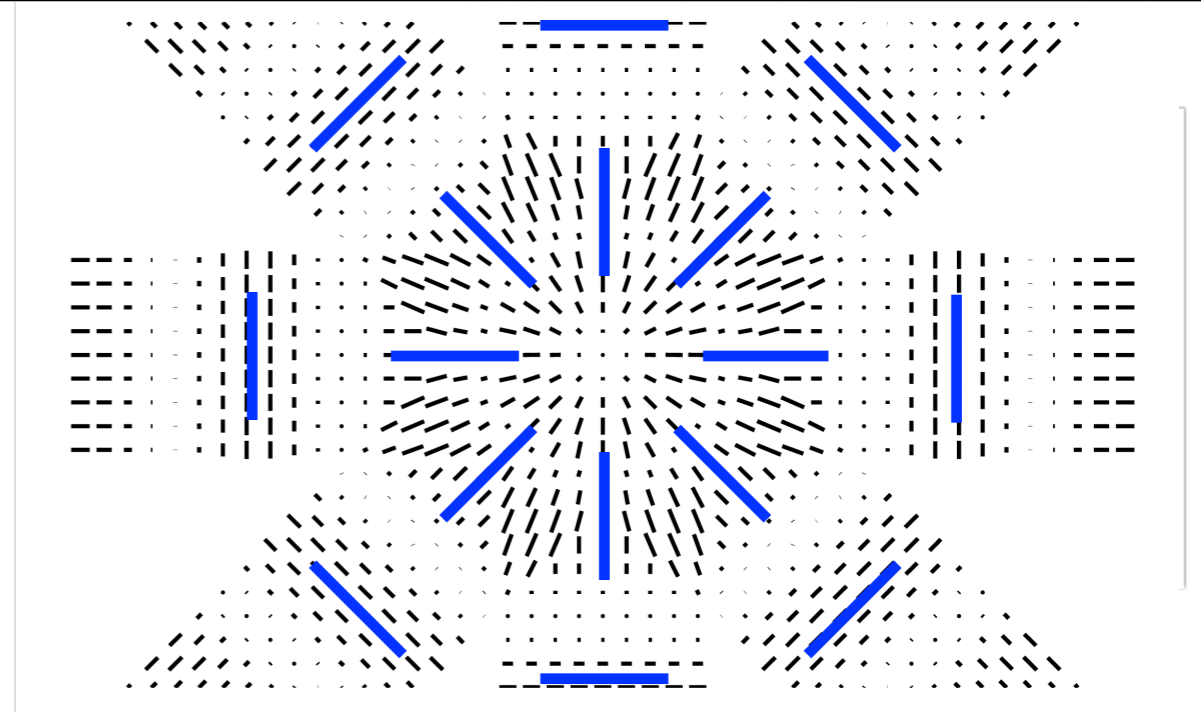
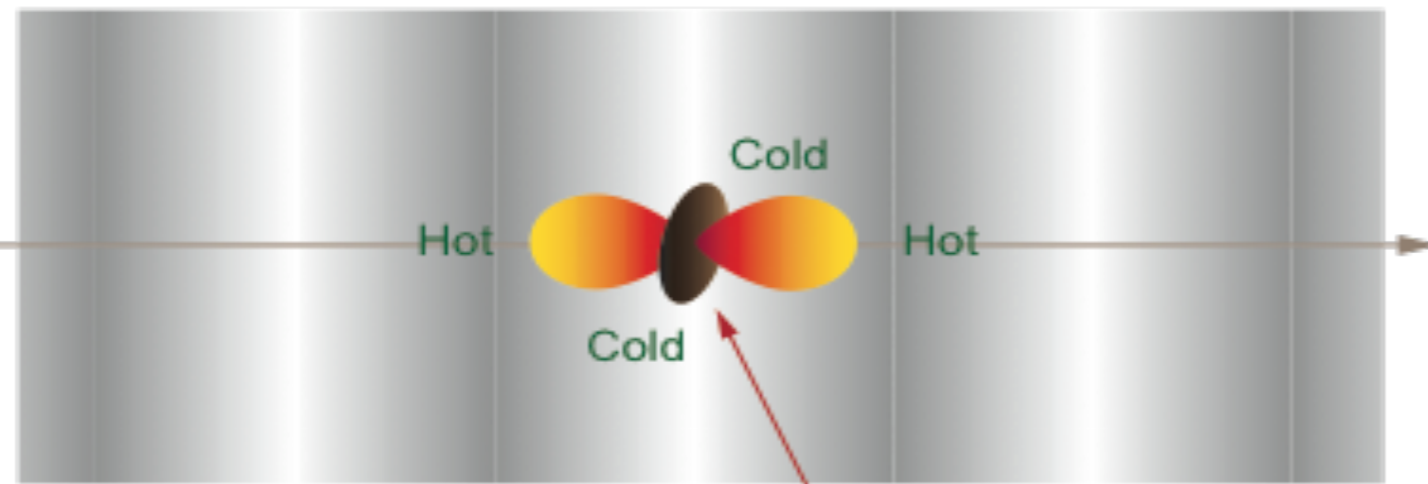
## B-mode Polarization





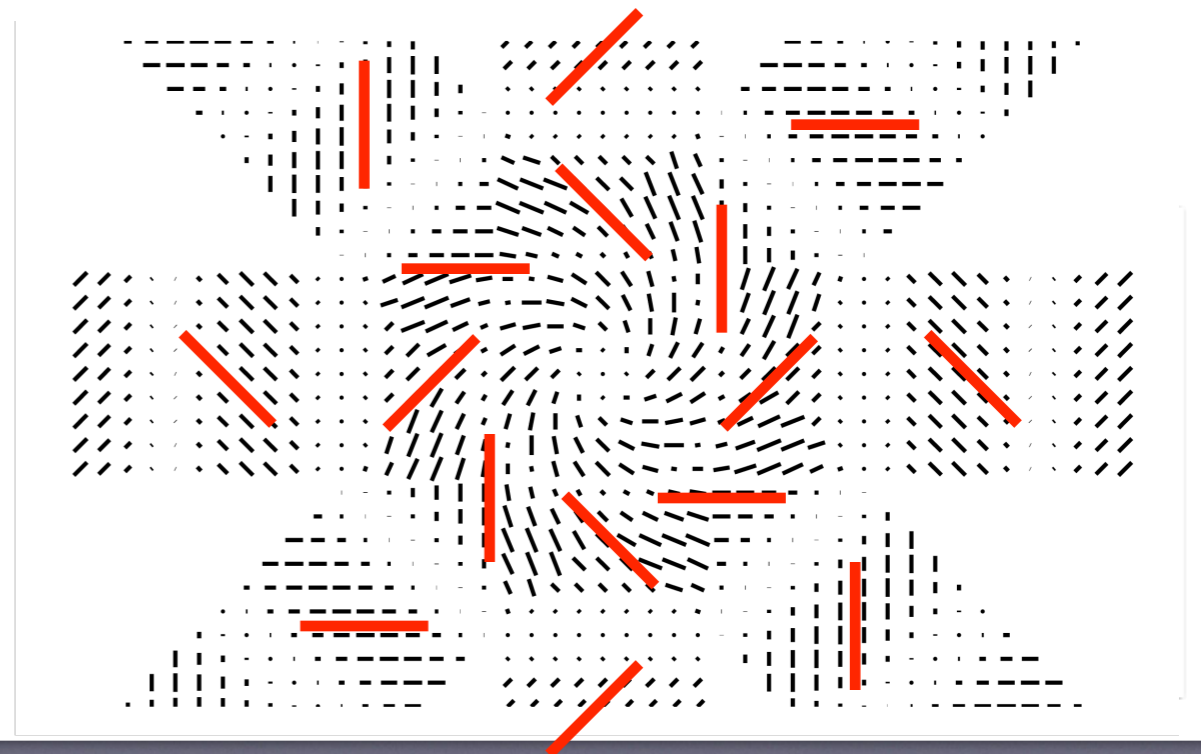
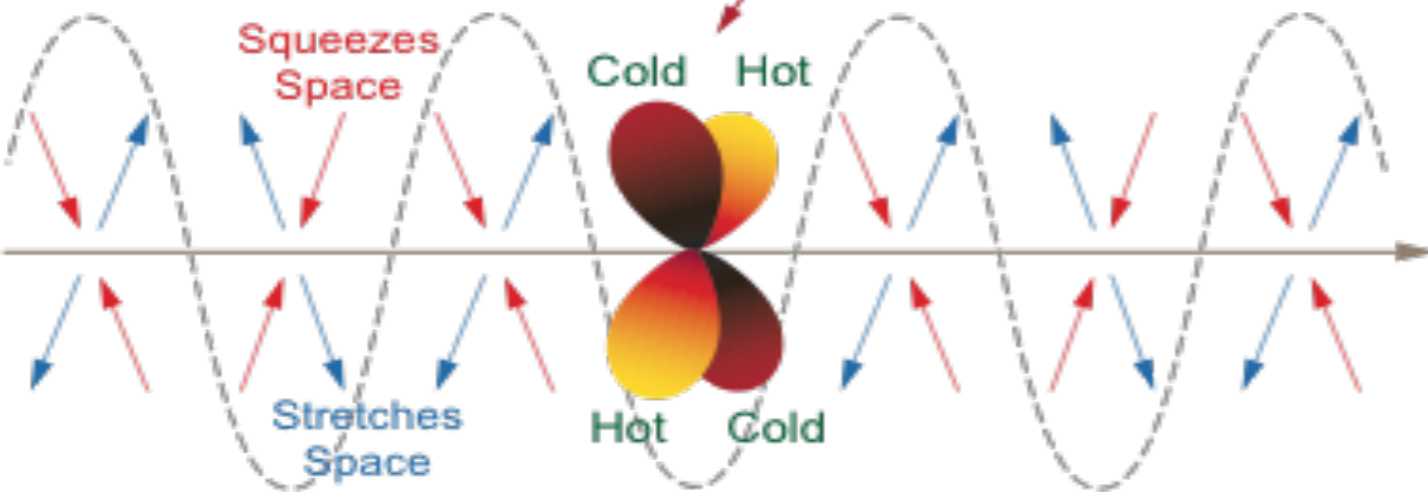
# Patterns of Polarization

Density Wave

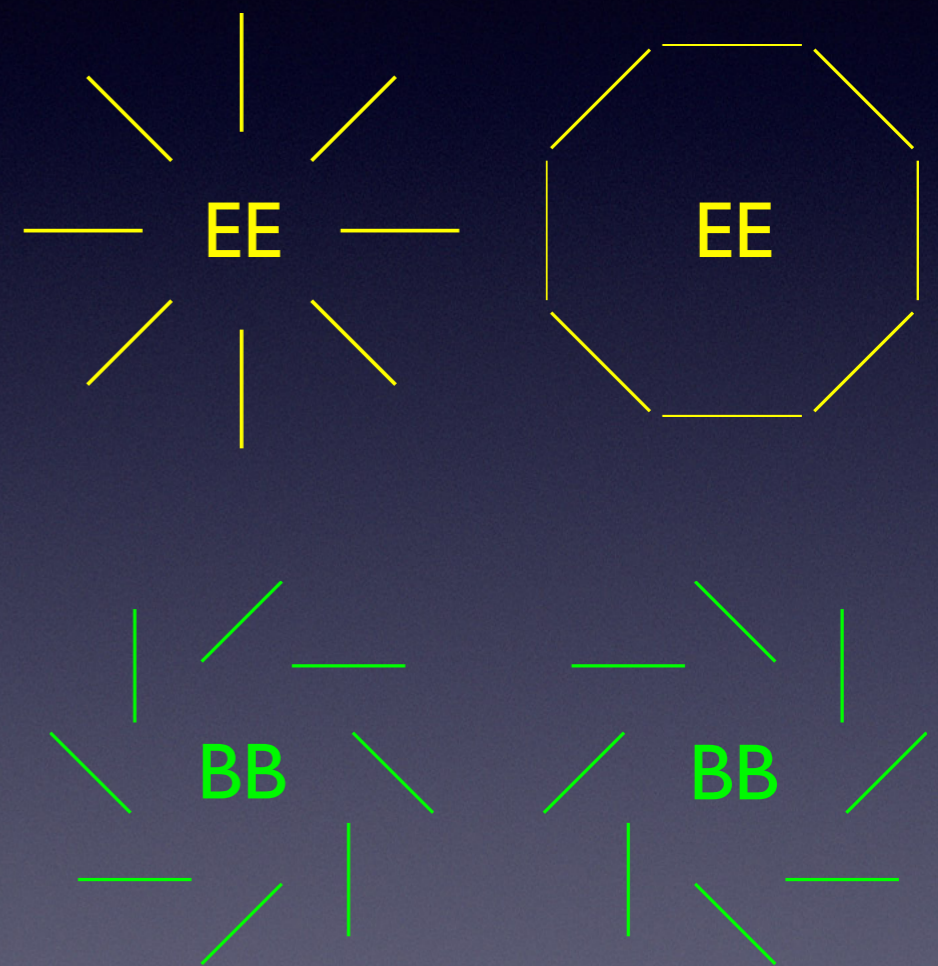
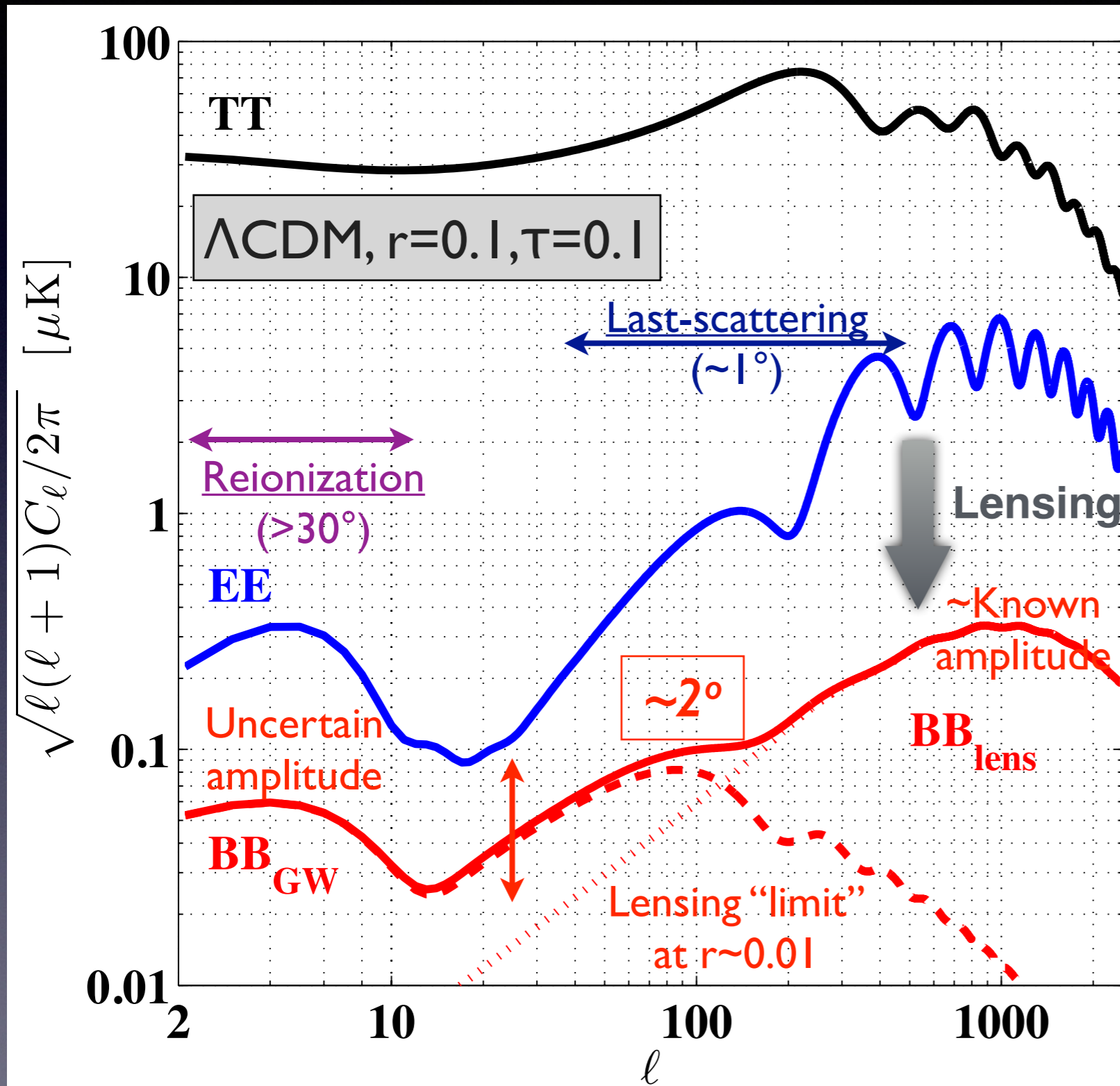


Temperature Pattern Seen by Electrons

Gravitational Wave



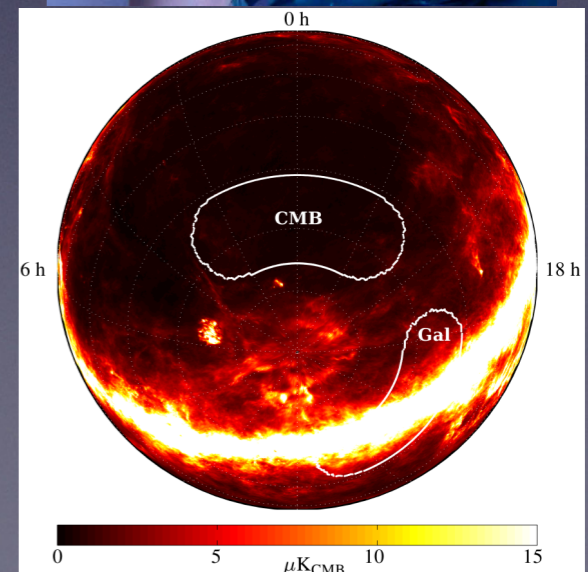
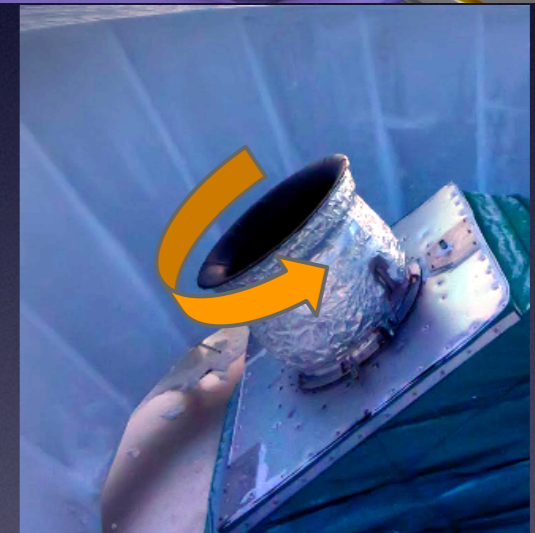
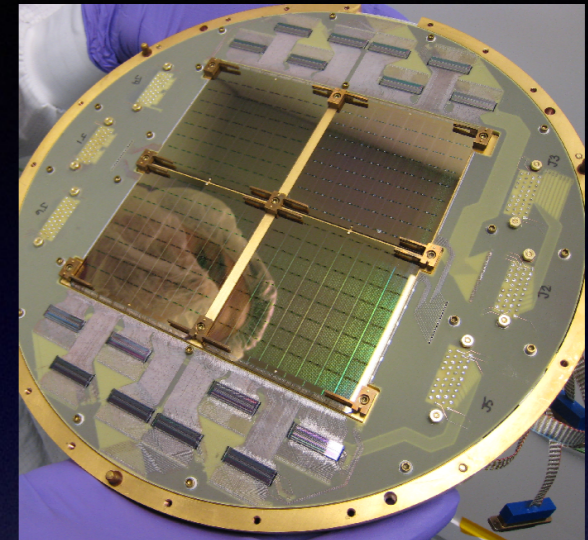
# The Lay of the Land



# The Challenge

Extremely faint signal demands a map that is...

- **Precise**  
Detectors approach photon noise limit  
Many detectors (*multiplexing*)
- **Accurate**  
Rigid control of polarized systematics
- **Uncontaminated**  
Avoidance (or subtraction) of polarized foregrounds
- *Not necessarily high angular resolution!*



# Inflation Investigators



**JPL**



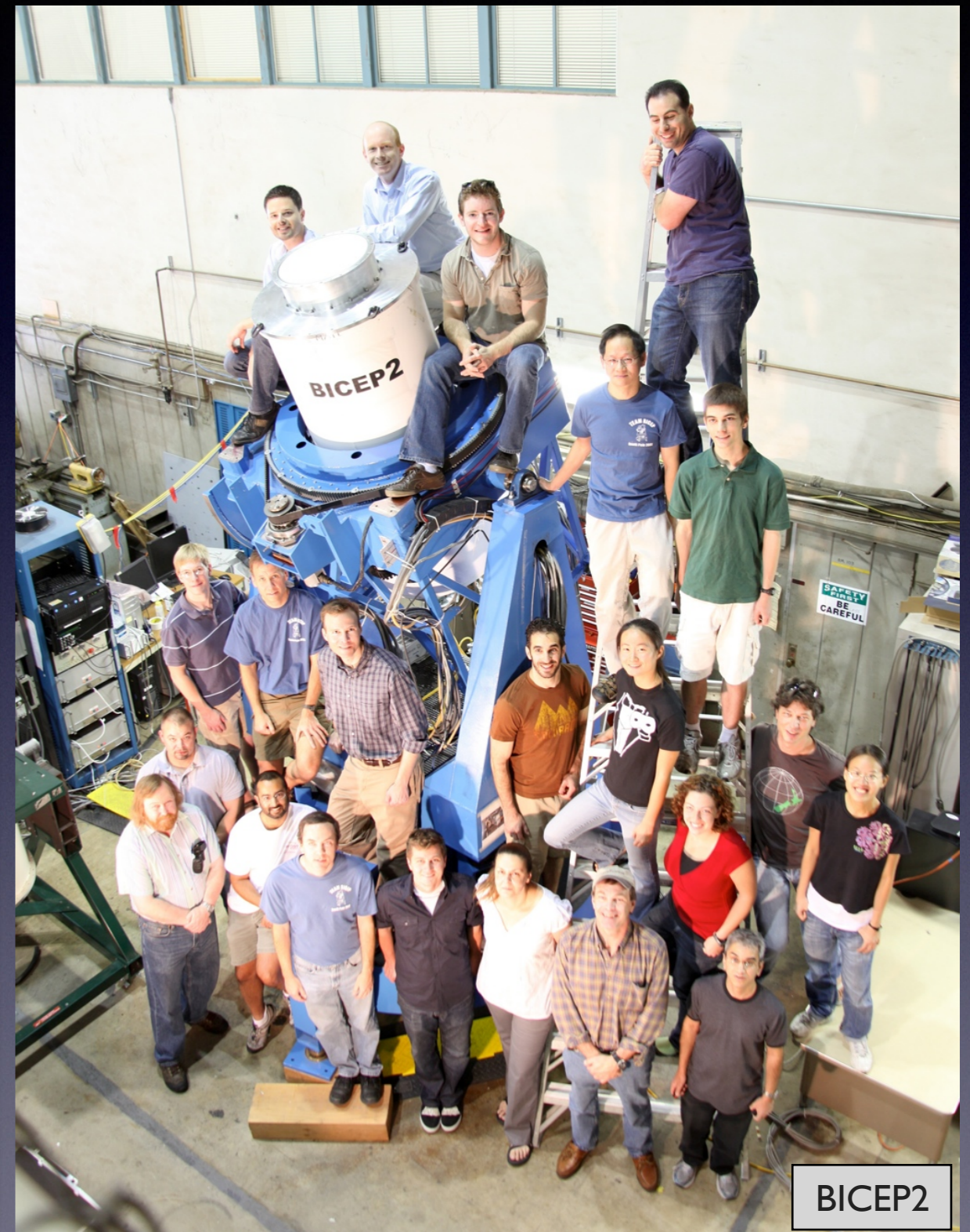
**NIST**

**CARDIFF**  
UNIVERSITY



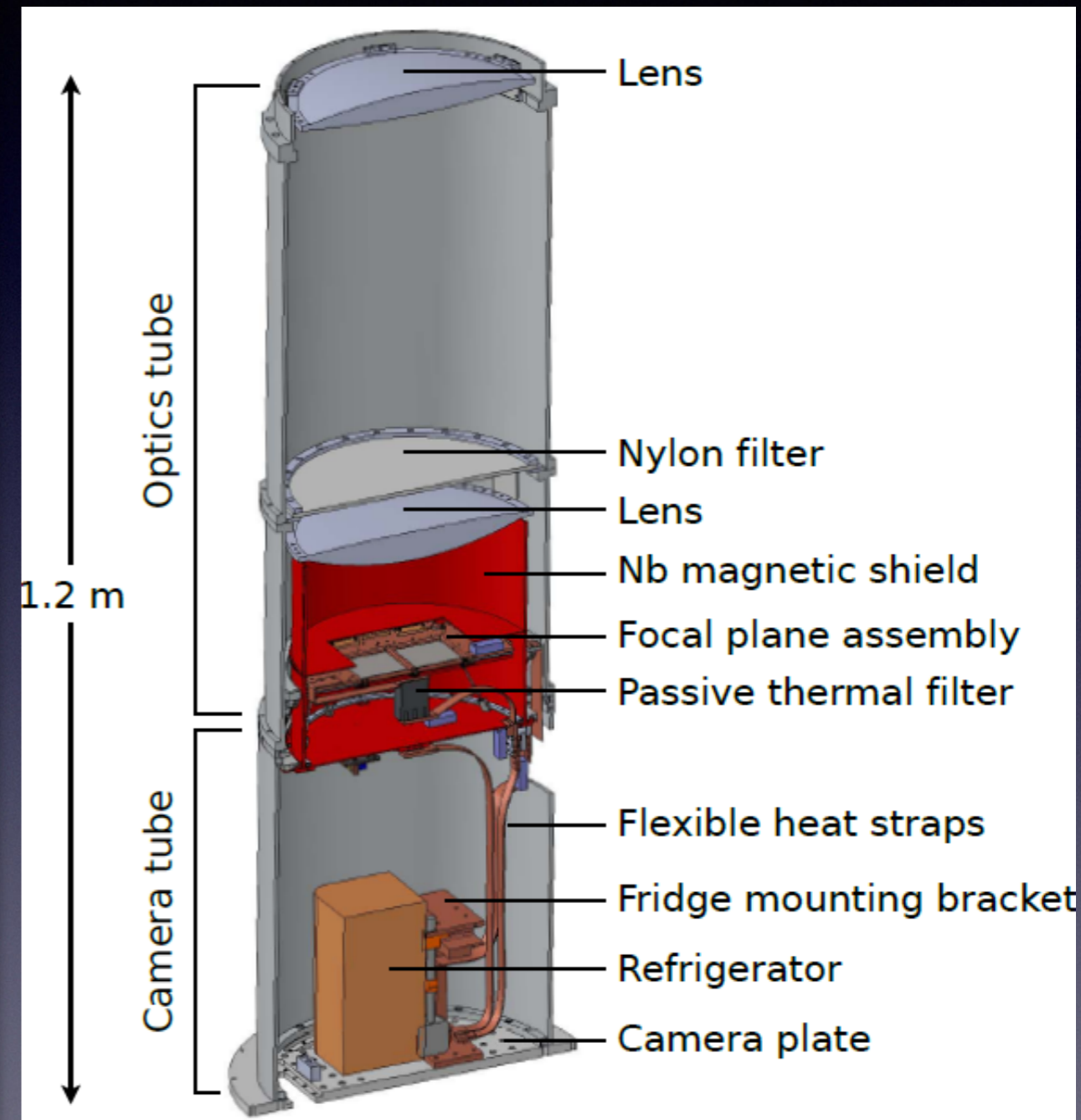
# A Targeted Strategy

- **Minimum aperture** (~26 cm) to resolve degree scales
- **Systematic control**
  - Cold (4.2K), on-axis optics
  - Bore sight rotation
- **Foreground avoidance**
  - Clean sky: ~400 sq. deg.
  - 150 GHz: low atmospheric and foreground emission
- **Deep mapping**
  - 3 years: 87 nK-deg!



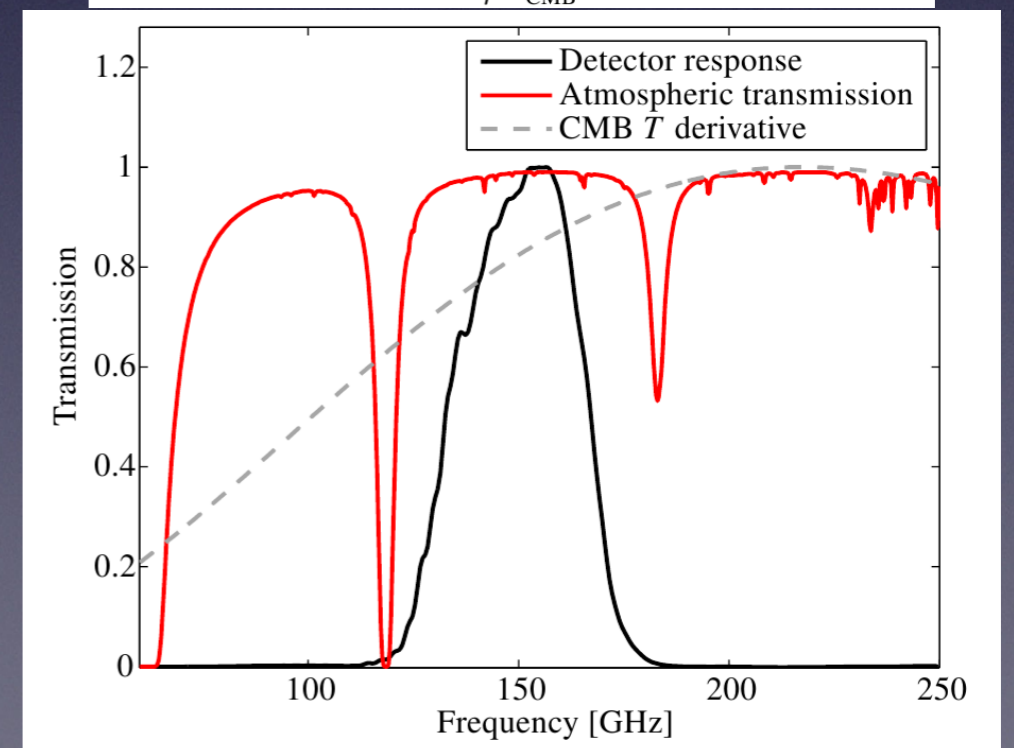
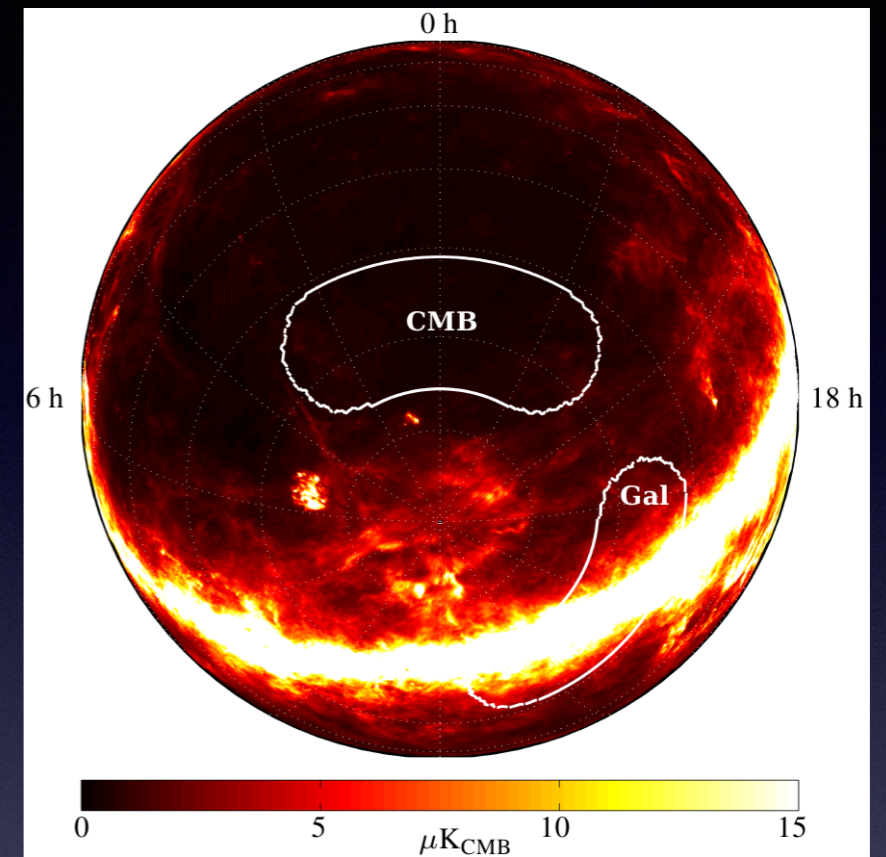
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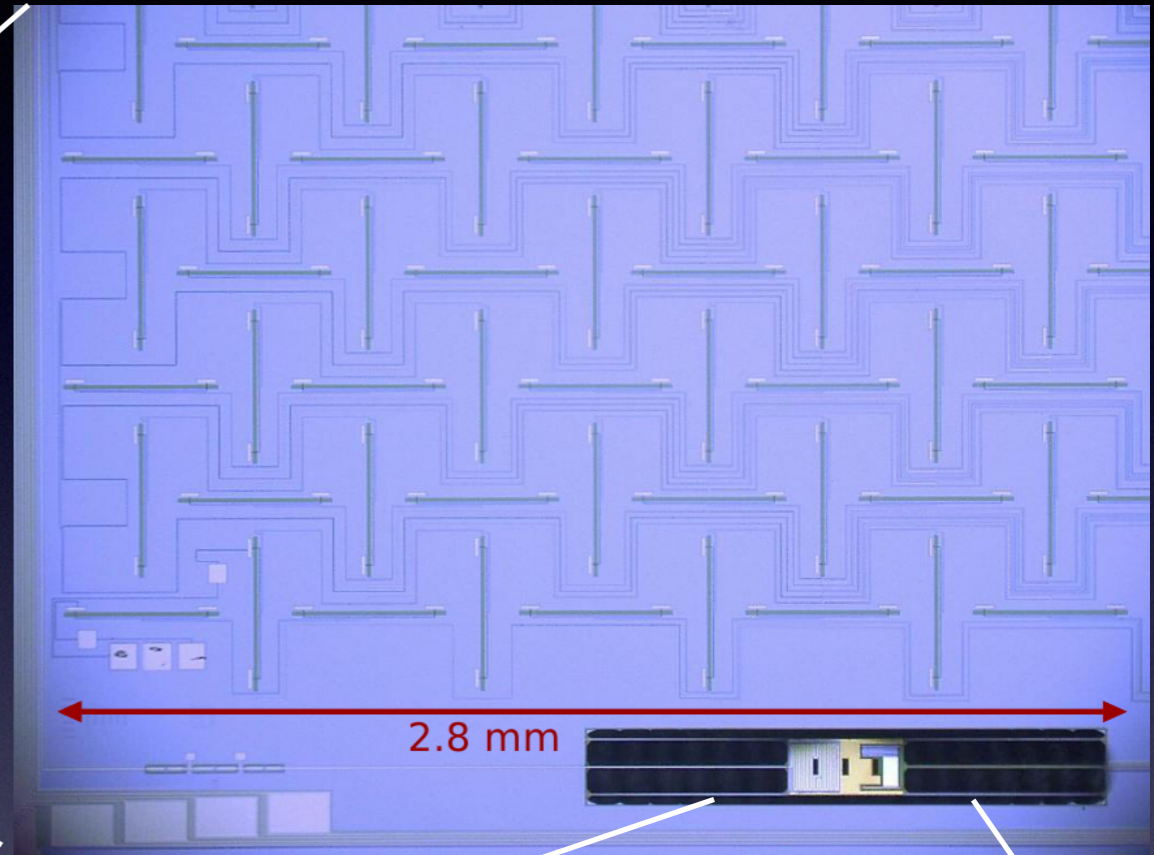
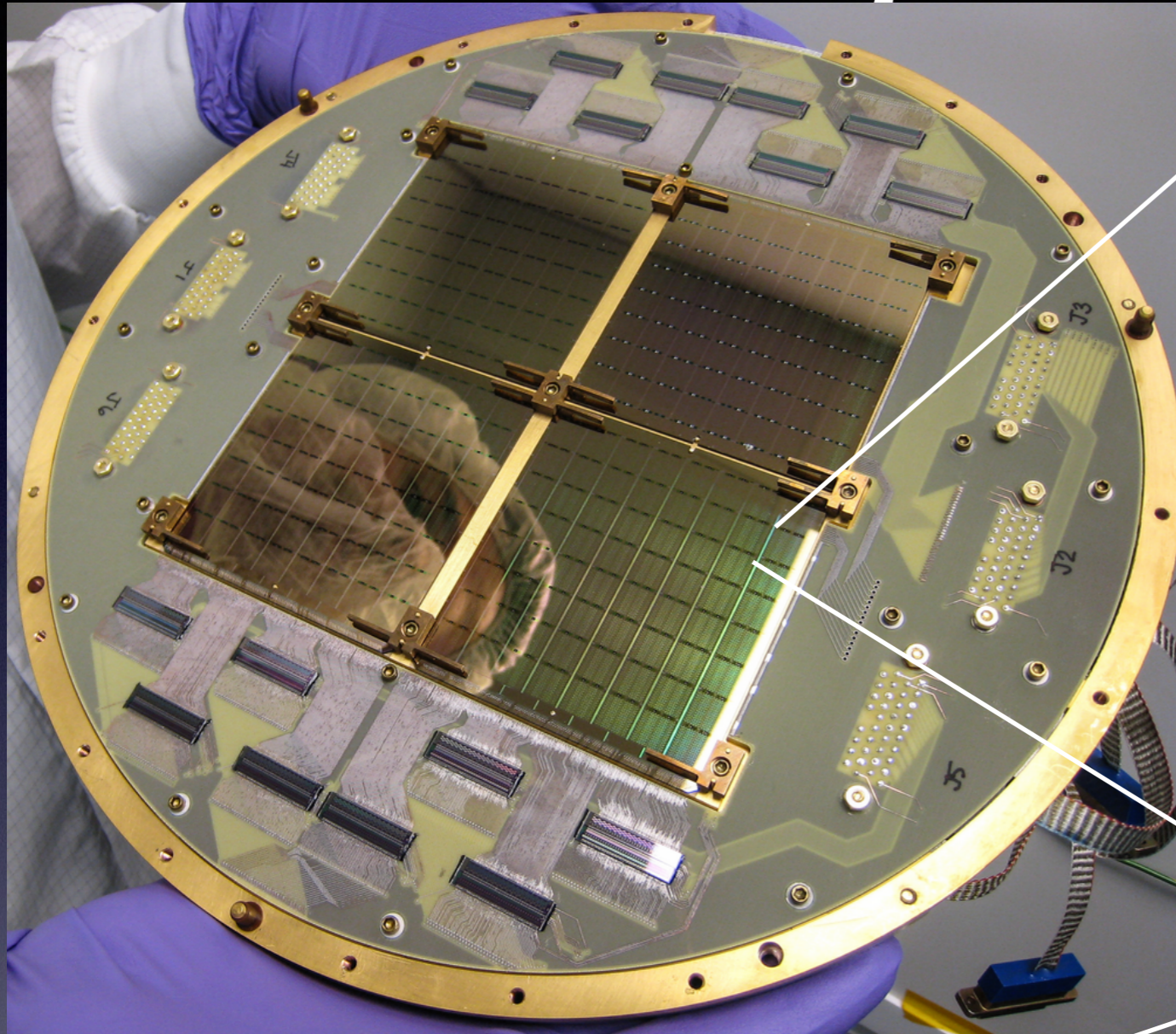


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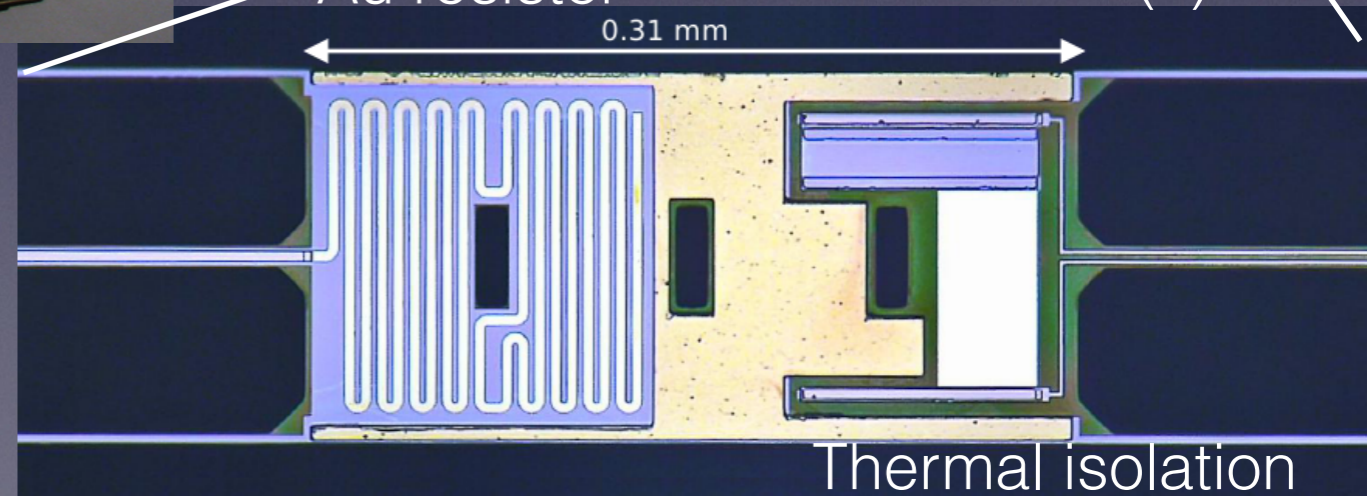
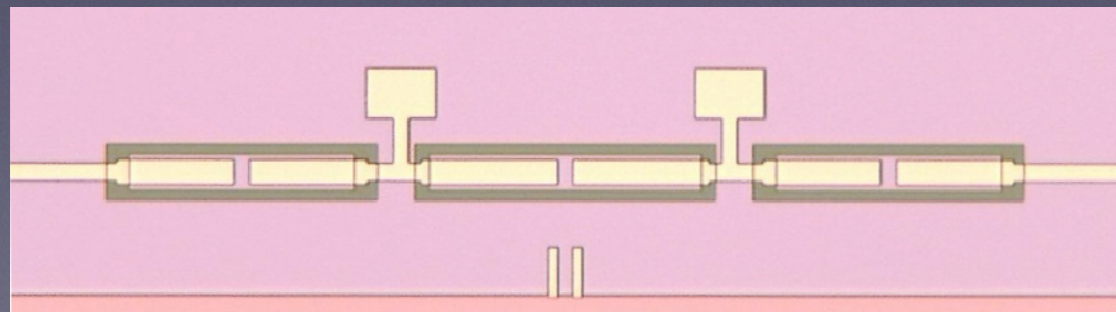


# Caltech/JPL Detectors



Au resistor

Superconducting thermometer(s)





# The View from the Bottom of the World

The Dark Sector Lab

BICEP

SPT

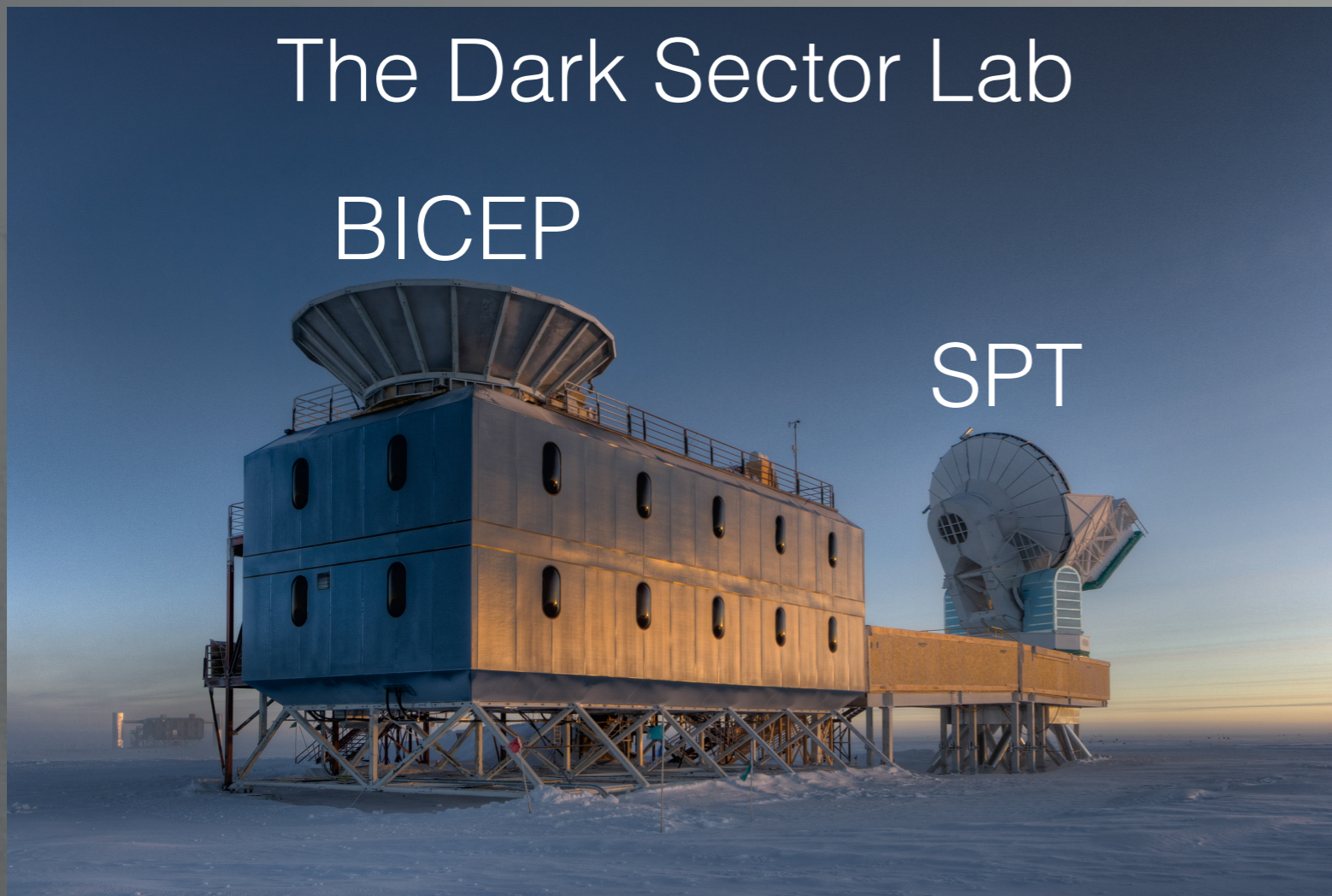
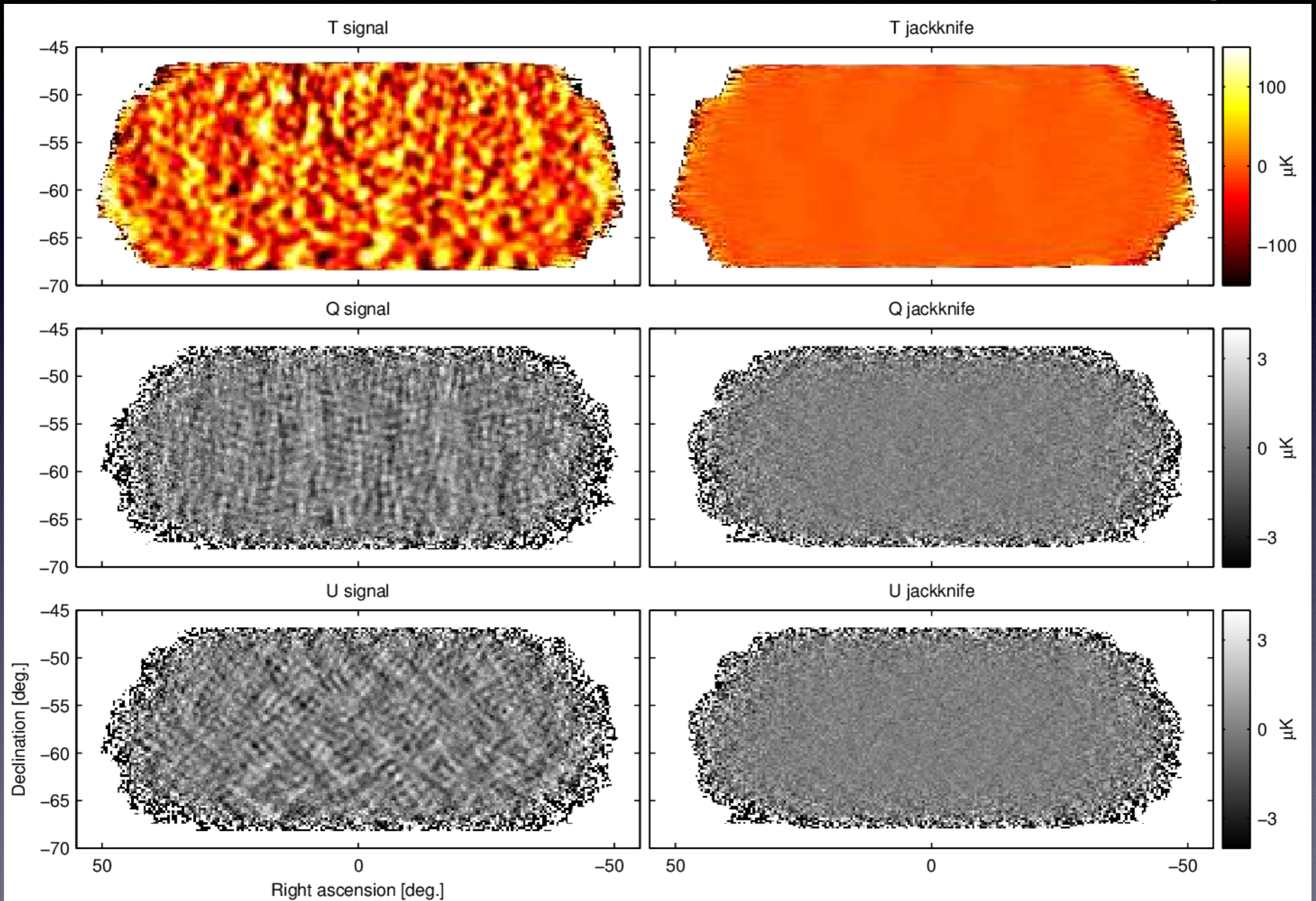




photo: Keith Vanderlinde

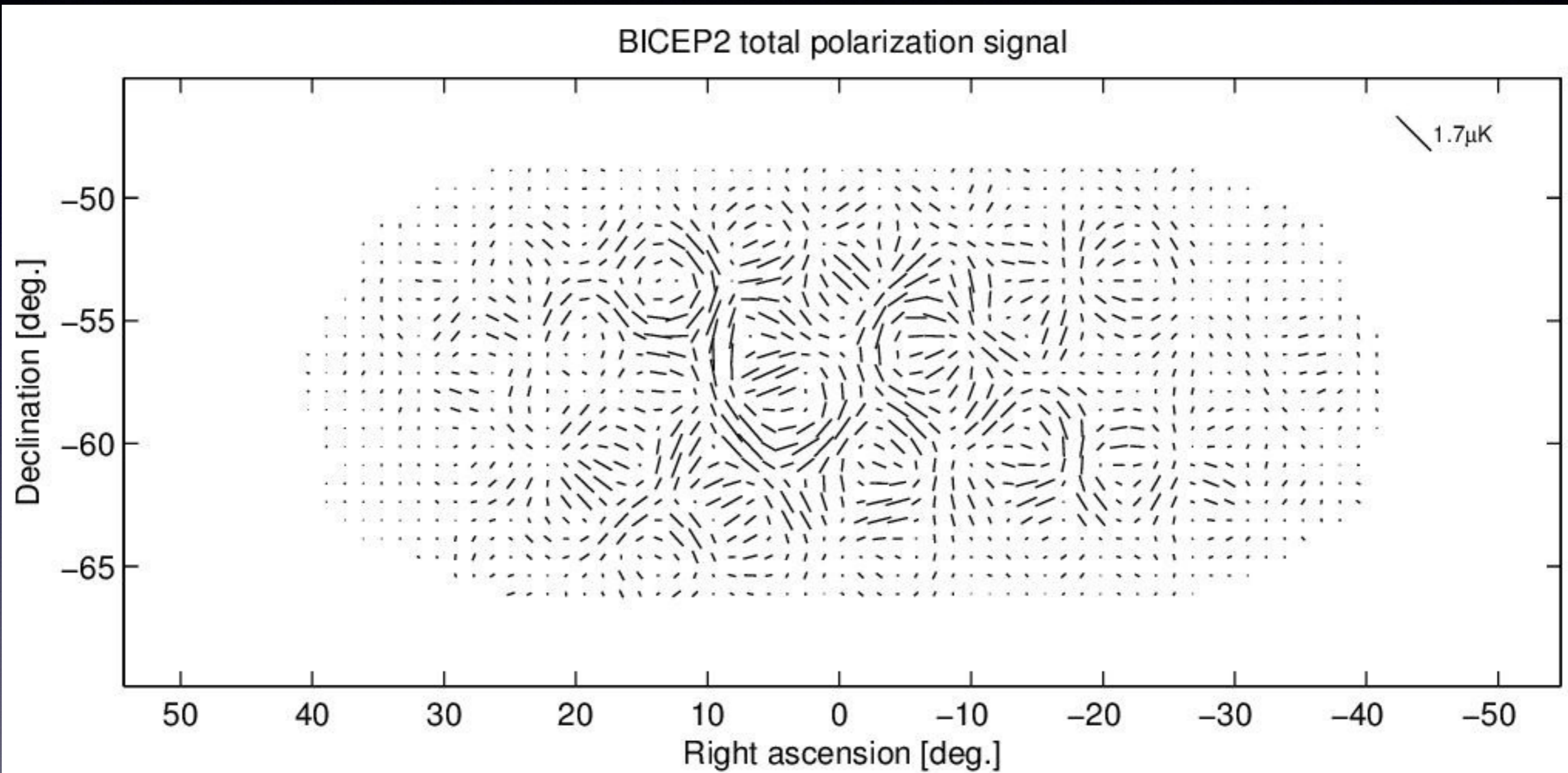
# BICEP2 T and Stokes Q/U Maps



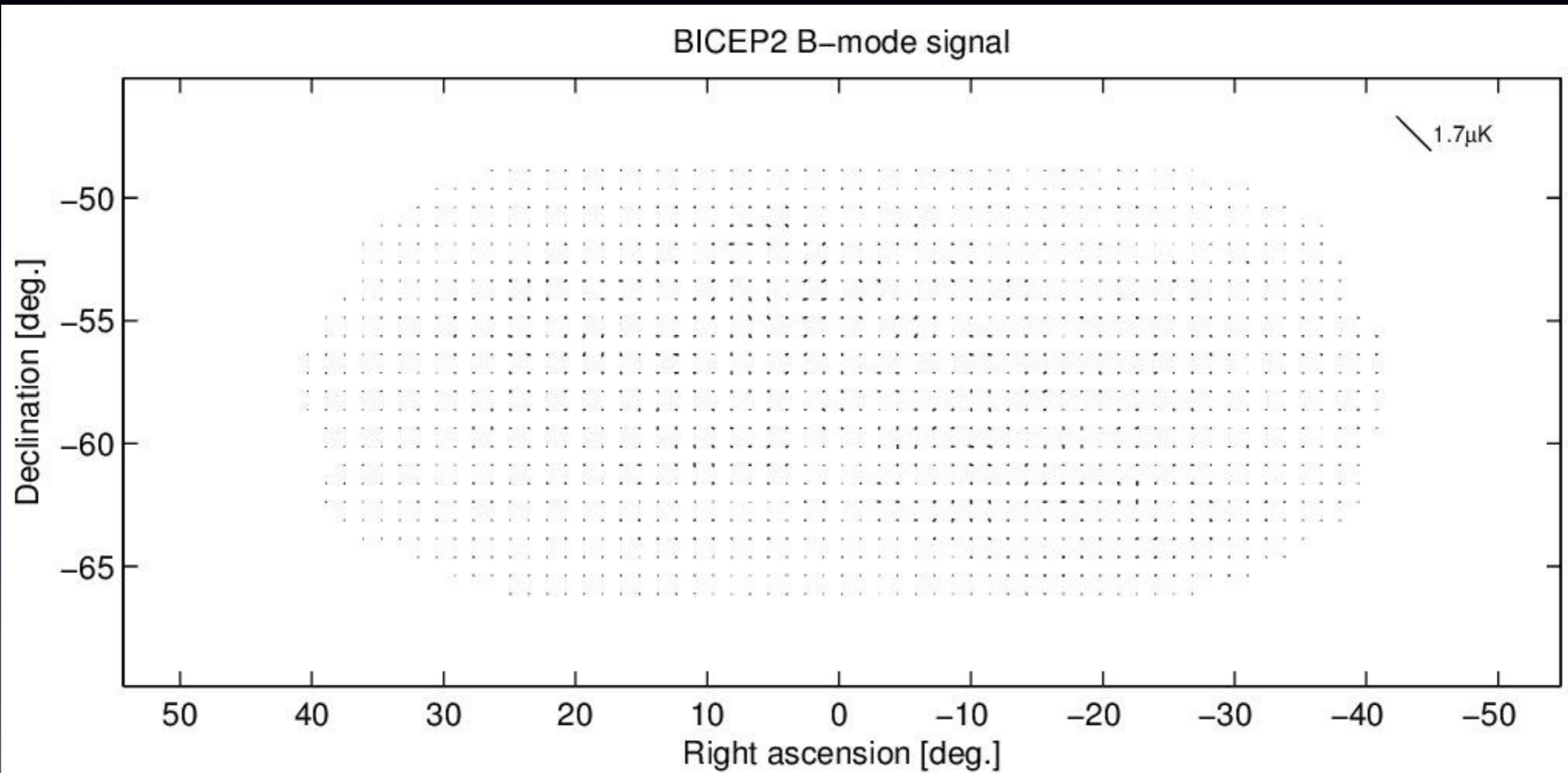
Sum maps

Difference maps

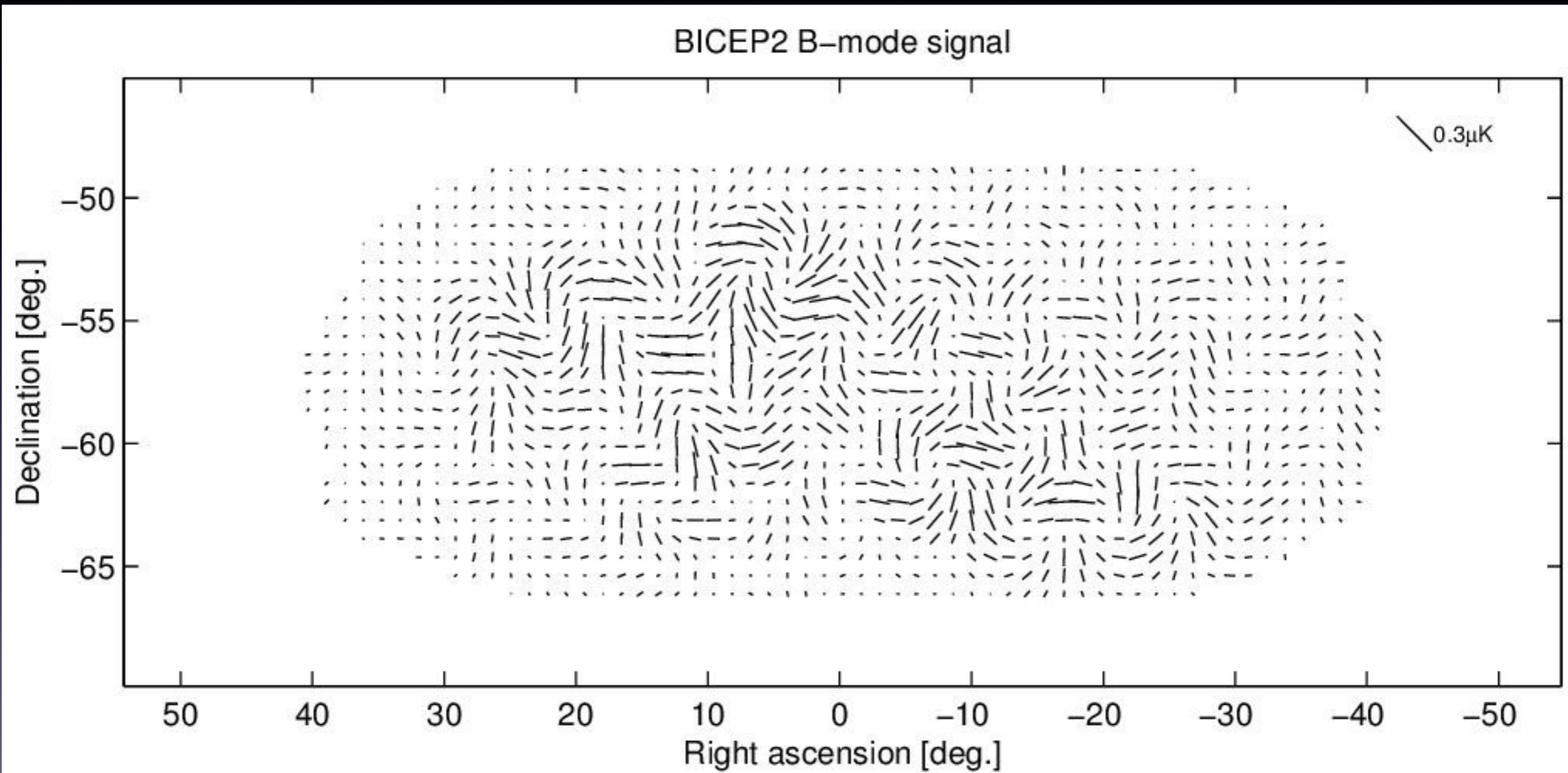
# Total Polarization



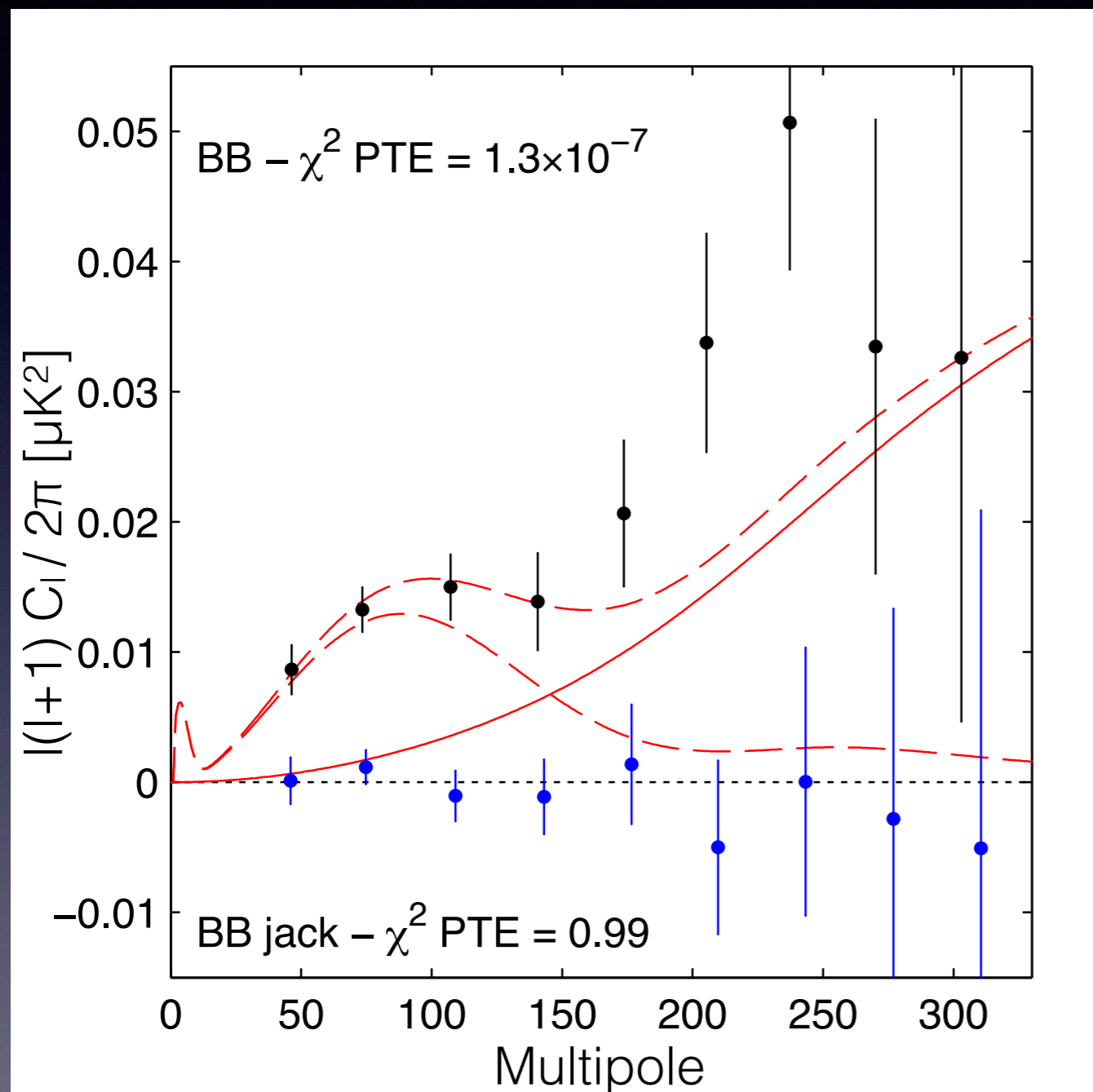
# B-mode Signal



# B-mode Signal



# BICEP2 B-mode Power Spectrum



- B-mode power spectrum
- temporal split jackknife
- lensed  $\Lambda\text{CDM}$
- -  $r=0.2$

- Clear excess of B-modes at low  $l$  above lensed- $\Lambda\text{CDM}$ 
  - PTE:  $1.3 \times 10^{-7}$
  - Significance:  $5.3\sigma$
- Good fit to expected inflationary signal spectrum

# What could this be?

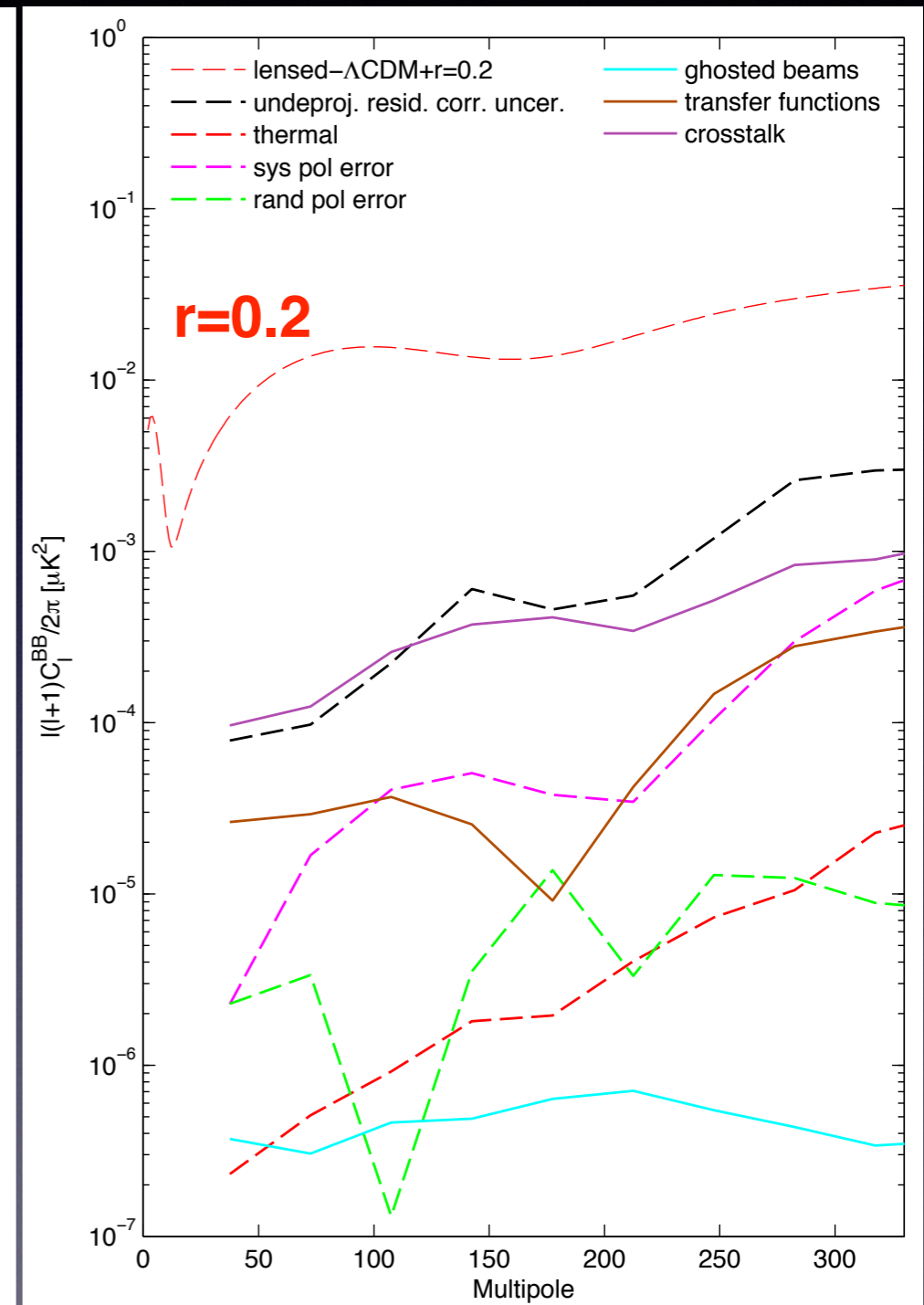
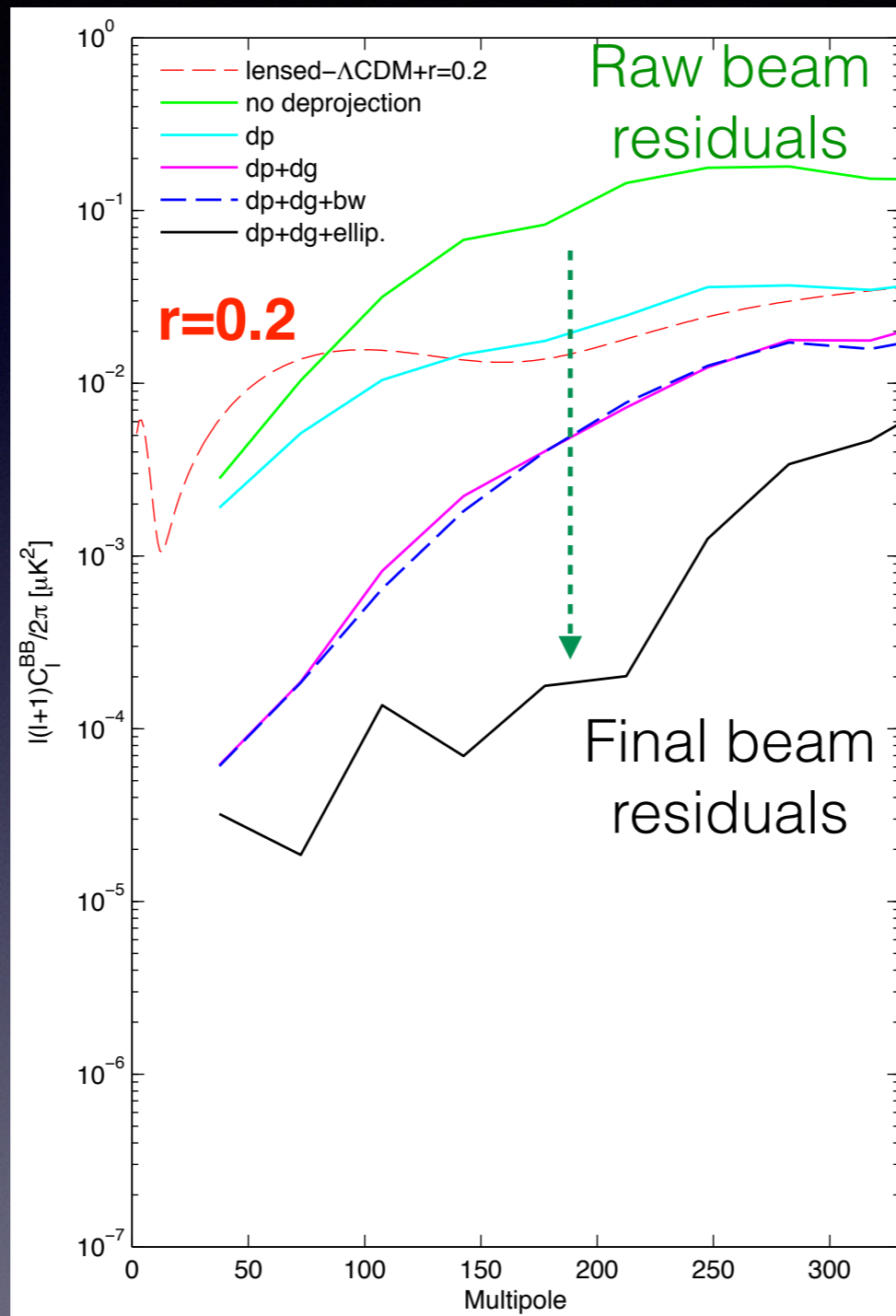
- **Instrumental systematics?**
- Galactic foregrounds?
- Cosmology?



# Systematic Errors

Simulate effect of **measured** beam and instrument imperfections.

We find with high confidence that the apparent signal *cannot be explained* by instrumental systematics



# Jackknife tests

TABLE 1  
JACKKNIFE PTE VALUES FROM  $\chi^2$  AND  $\chi$  (SUM-OF-DEVIATION)  
TESTS

Jackknife	Bandpowers 1-5 $\chi^2$	Bandpowers 1-9 $\chi^2$	Bandpowers 1-5 $\chi$	Bandpowers 1-9 $\chi$
<b>Deck jackknife</b>				
EE	0.046	0.030	0.164	0.299
BB	0.774	0.329	0.240	0.082
EB	0.337	0.643	0.204	0.267
<b>Scan Dir jackknife</b>				
EE	0.483	0.762	0.978	0.938
BB	0.531	0.573	0.896	0.551
EB	0.898	0.806	0.725	0.890
<b>Tag Split jackknife</b>				
EE	0.541	0.377	0.916	0.938
BB	0.902	0.992	0.449	0.585
EB	0.477	0.689	0.856	0.615
<b>Tile jackknife</b>				
EE	0.004	0.010	0.000	0.002
BB	0.794	0.752	0.565	0.331
EB	0.172	0.419	0.962	0.790
<b>Phase jackknife</b>				
EE	0.673	0.409	0.126	0.339
BB	0.591	0.739	0.842	0.944
EB	0.529	0.577	0.840	0.659
<b>Mux Col jackknife</b>				
EE	0.812	0.587	0.196	0.204
BB	0.826	0.972	0.293	0.283
EB	0.866	0.968	0.876	0.697
<b>Alt Deck jackknife</b>				
EE	0.004	0.004	0.070	0.236
BB	0.397	0.176	0.381	0.086
EB	0.150	0.060	0.170	0.291
<b>Mux Row jackknife</b>				
EE	0.052	0.178	0.653	0.739
BB	0.345	0.361	0.032	0.008
EB	0.529	0.226	0.024	0.048
<b>Tile/Deck jackknife</b>				
EE	0.048	0.088	0.144	0.132
BB	0.908	0.840	0.629	0.269
EB	0.050	0.154	0.591	0.591
<b>Focal Plane inner/outer jackknife</b>				
EE	0.230	0.597	0.022	0.090
BB	0.216	0.531	0.046	0.092
EB	0.036	0.042	0.850	0.838
<b>Tile top/bottom jackknife</b>				
EE	0.289	0.347	0.459	0.599
BB	0.293	0.236	0.154	0.028
EB	0.545	0.683	0.902	0.932
<b>Tile inner/outer jackknife</b>				
EE	0.727	0.533	0.128	0.485
BB	0.255	0.086	0.421	0.036
EB	0.465	0.737	0.208	0.168
<b>Moon jackknife</b>				
EE	0.499	0.689	0.481	0.679
BB	0.144	0.287	0.898	0.858
EB	0.289	0.359	0.531	0.307
<b>A/B offset best/worst</b>				
EE	0.317	0.311	0.868	0.709
BB	0.114	0.064	0.307	0.094
EB	0.589	0.872	0.599	0.790

## Splits by boresight rotation

Amplifies differential pointing in comparison to fully added data. Check of deprojection.

## Splits by time

Checks for contamination on long (“Tag Split”) and short (“Scan Dir”) timescales. Short timescales probe detector transfer functions.

## Splits by channel selection

Checks for contamination in channel subgroups, divided by focal plane location, tile location, and readout electronics grouping

## Splits by possible external contamination

Checks for contamination from ground-fixed signals, such as polarized sky or magnetic fields, or the moon

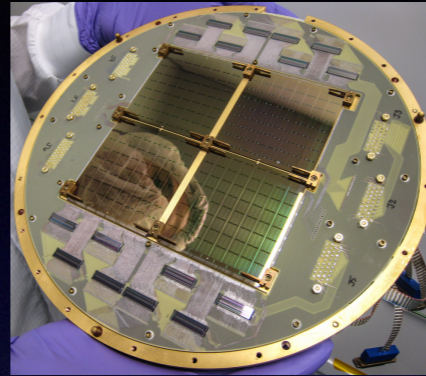
## Splits to check intrinsic detector properties

Checks for contamination from detectors with best/worst differential pointing. “Tile/dk” divides the data by the orientation of the detector on the sky.

# Cross Correlations



BICEP1: Feedhorns, NTDs,  
150 and 100 GHz

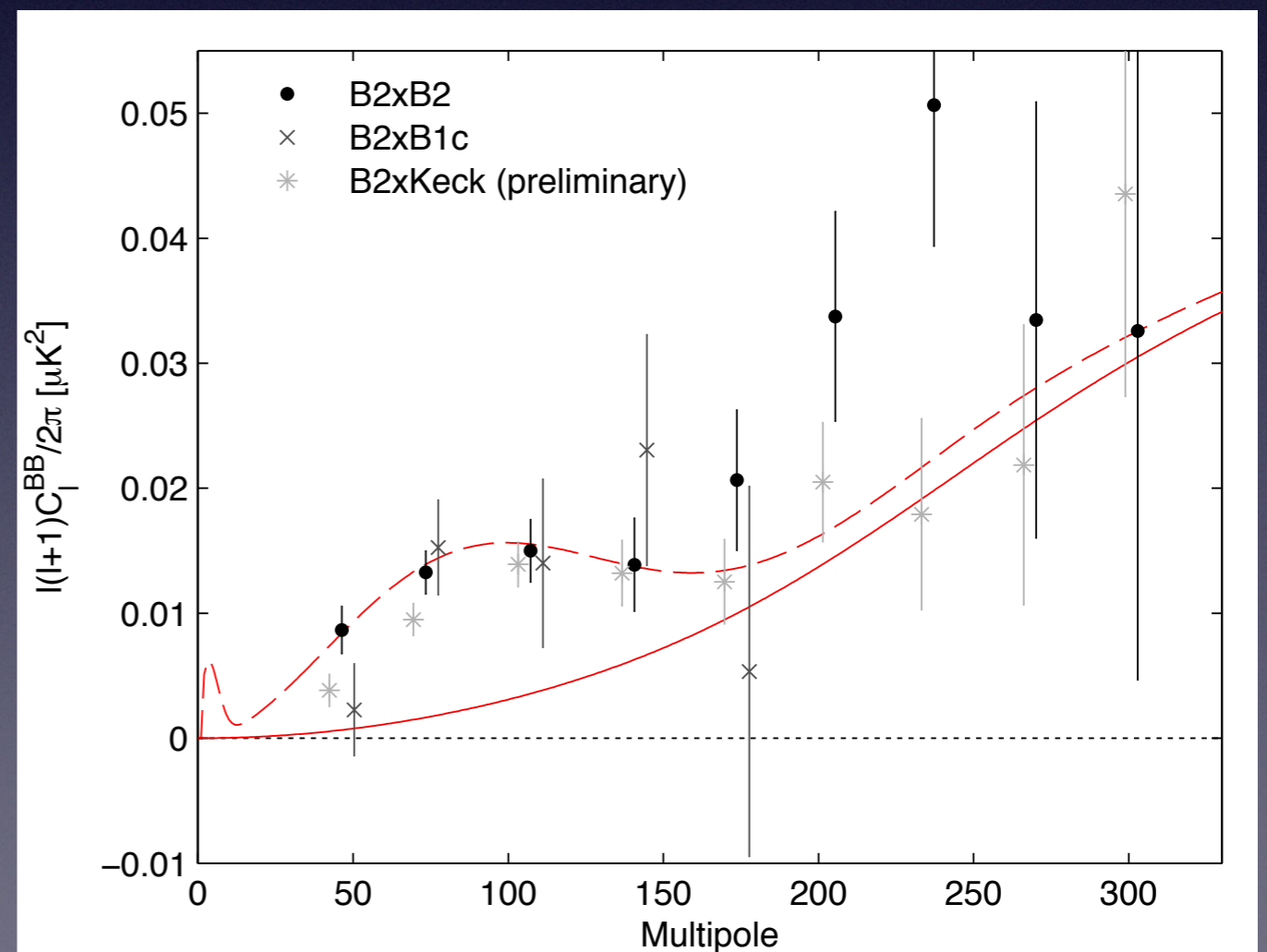


BICEP2: Phased antenna  
array, TESSs, 150 GHz



Keck Array: Phased antenna  
array, TESSs, 150 GHz

- **3.5 $\sigma$**  detection of BB in cross with color-combined BICEP1 ( $r_{\text{max}}=0.19$ )
- Excess power also evident in cross with 2 years of **Keck Array** data (*150 GHz, preliminary*)!



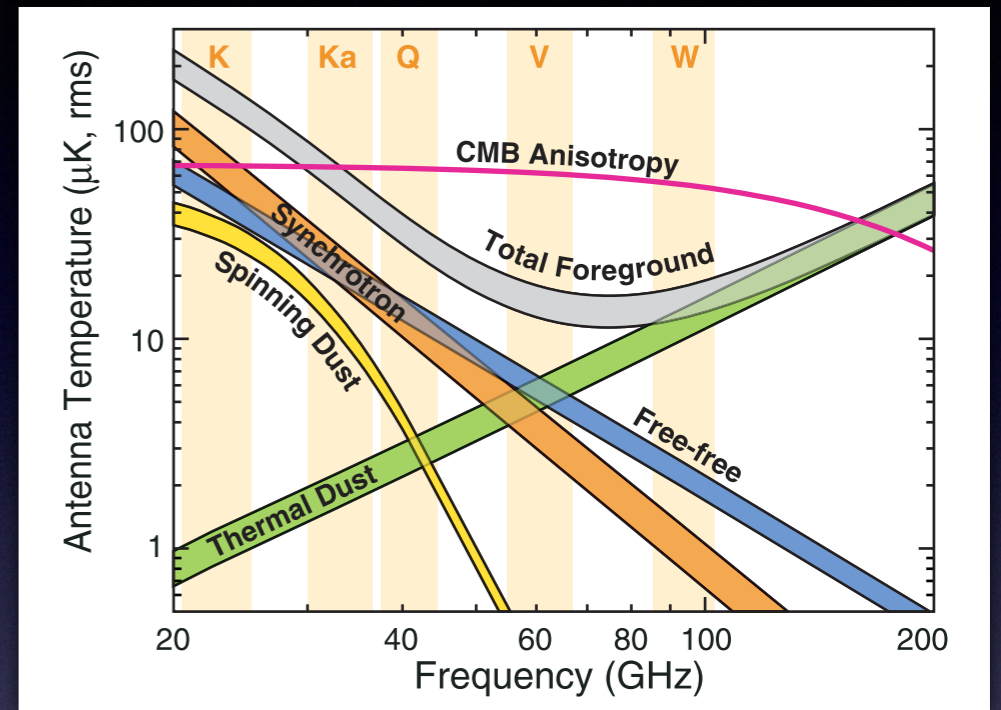
# What could this be?

- Instrumental systematics?
- **Galactic foregrounds?**
- Cosmology?

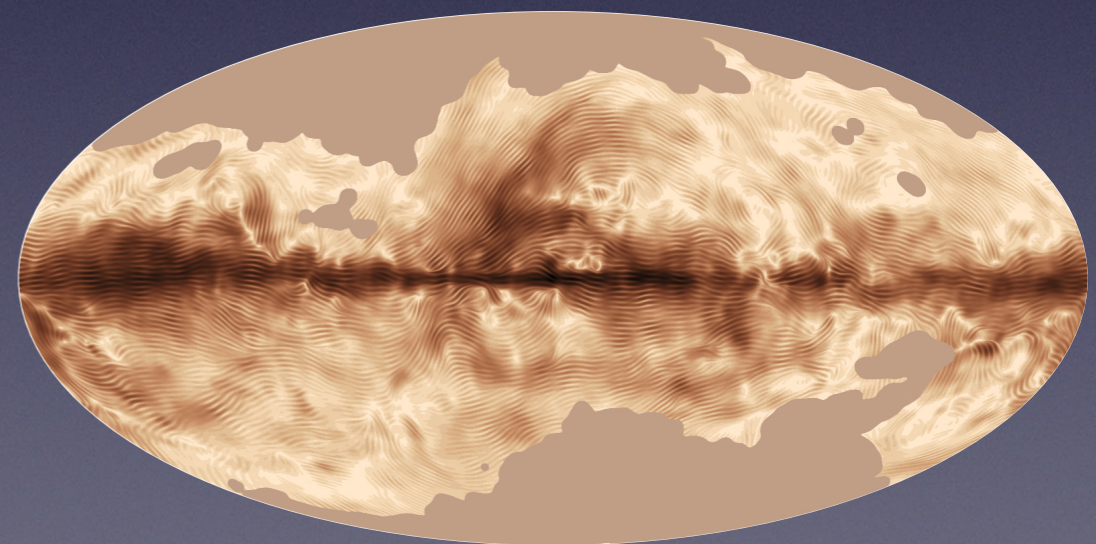
# Polarized Foregrounds

**Any** polarized astrophysical emission between last scattering and us!

- **Synchrotron** “Red”:  $\sim \nu^{-3}$
- **Dust** “Blue”:  $\sim \nu^{+1.75}$
- **Point sources**



WMAP *unpolarized*



Planck 2014 magnetic field

# Polarized Foregrounds

**Any** polarized astrophysical emission between last scattering and us!

- **Synchrotron** “Red”:  $\sim \nu^{-3}$

*No correlation with WMAP-K*

- **Dust** “Blue”:  $\sim \nu^{+1.75}$

*Brighter than existing models*

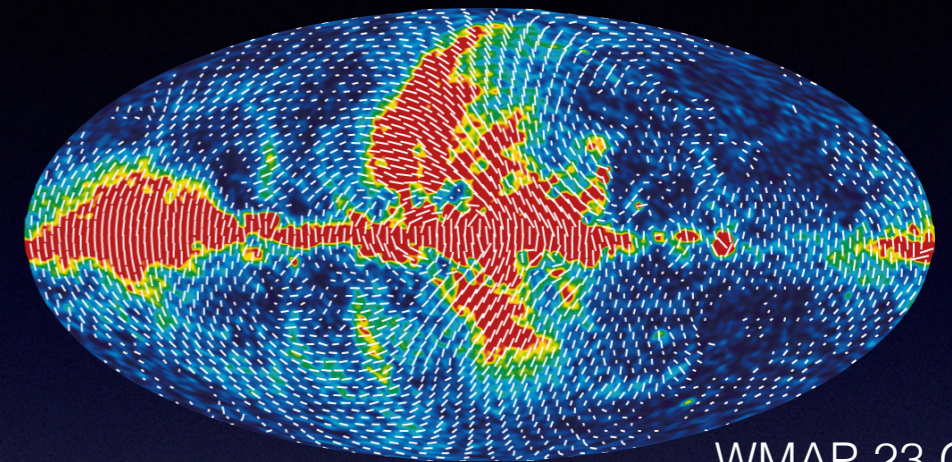
*Lack of cross-correlation*

*Angular spectrum consistent with GW signal*

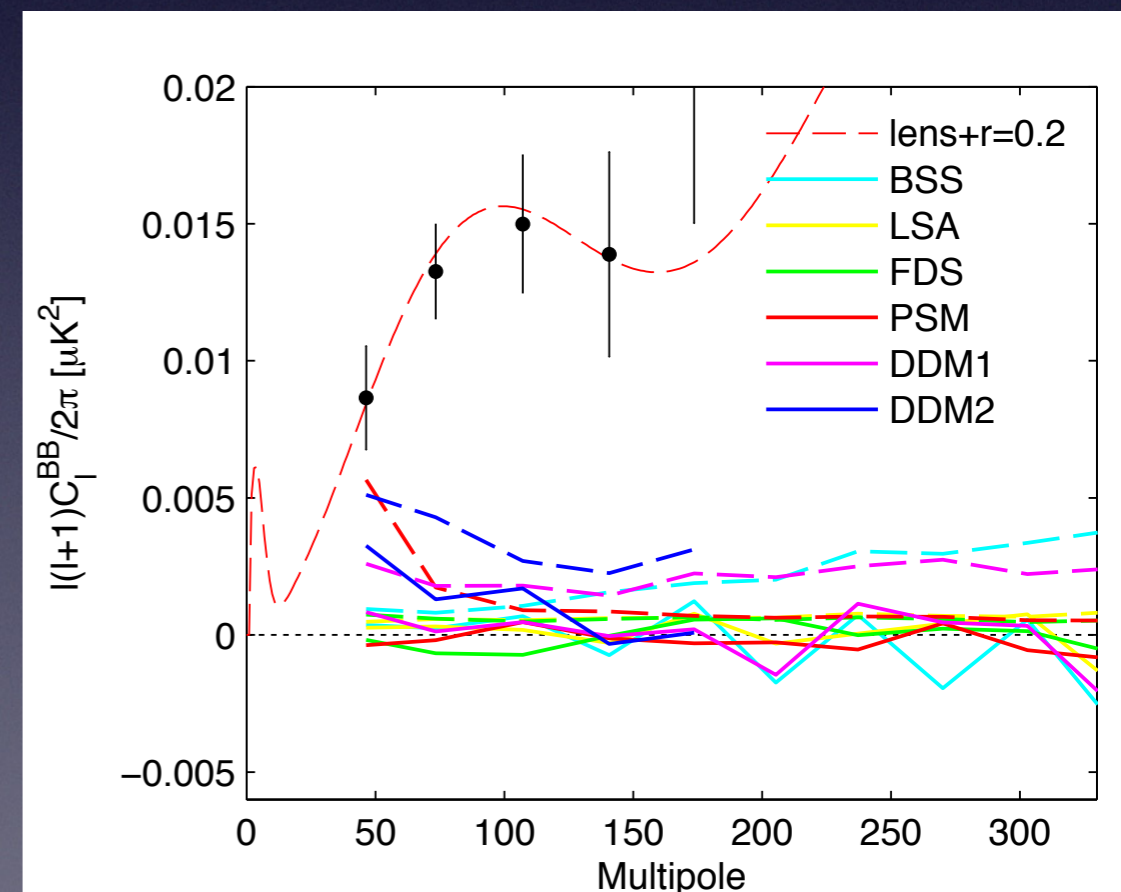
- **Point sources**

*No cross-correlation with source catalogs*

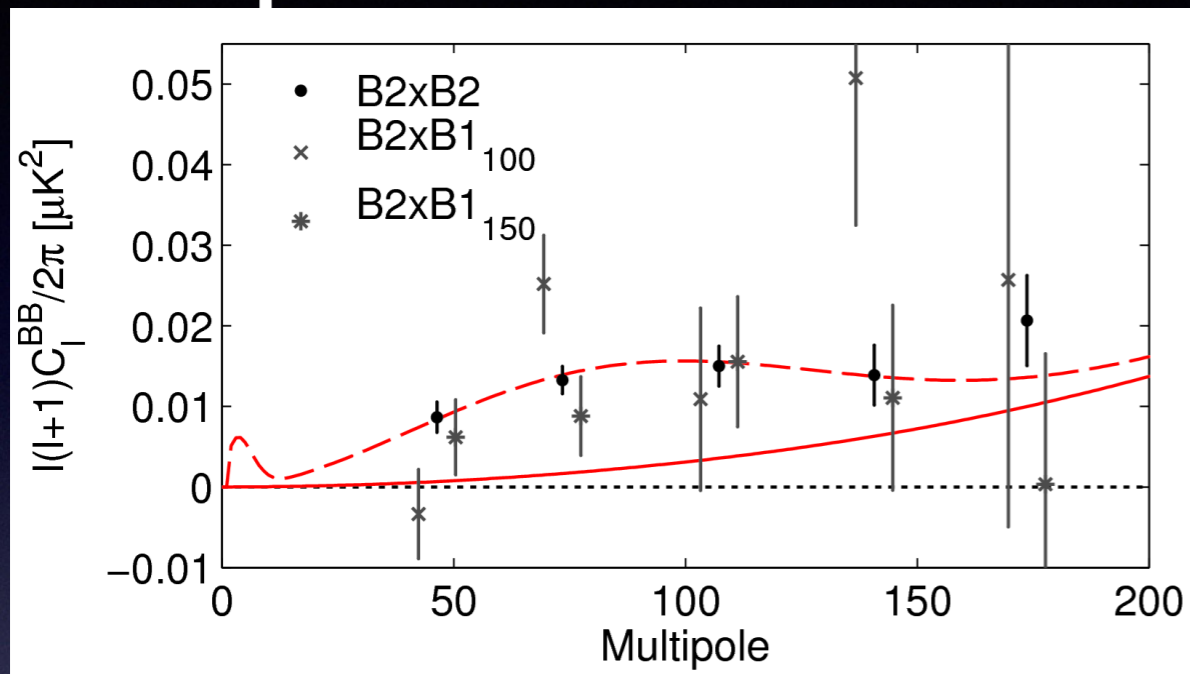
*(Planck, ATCA)*



WMAP 23 GHz

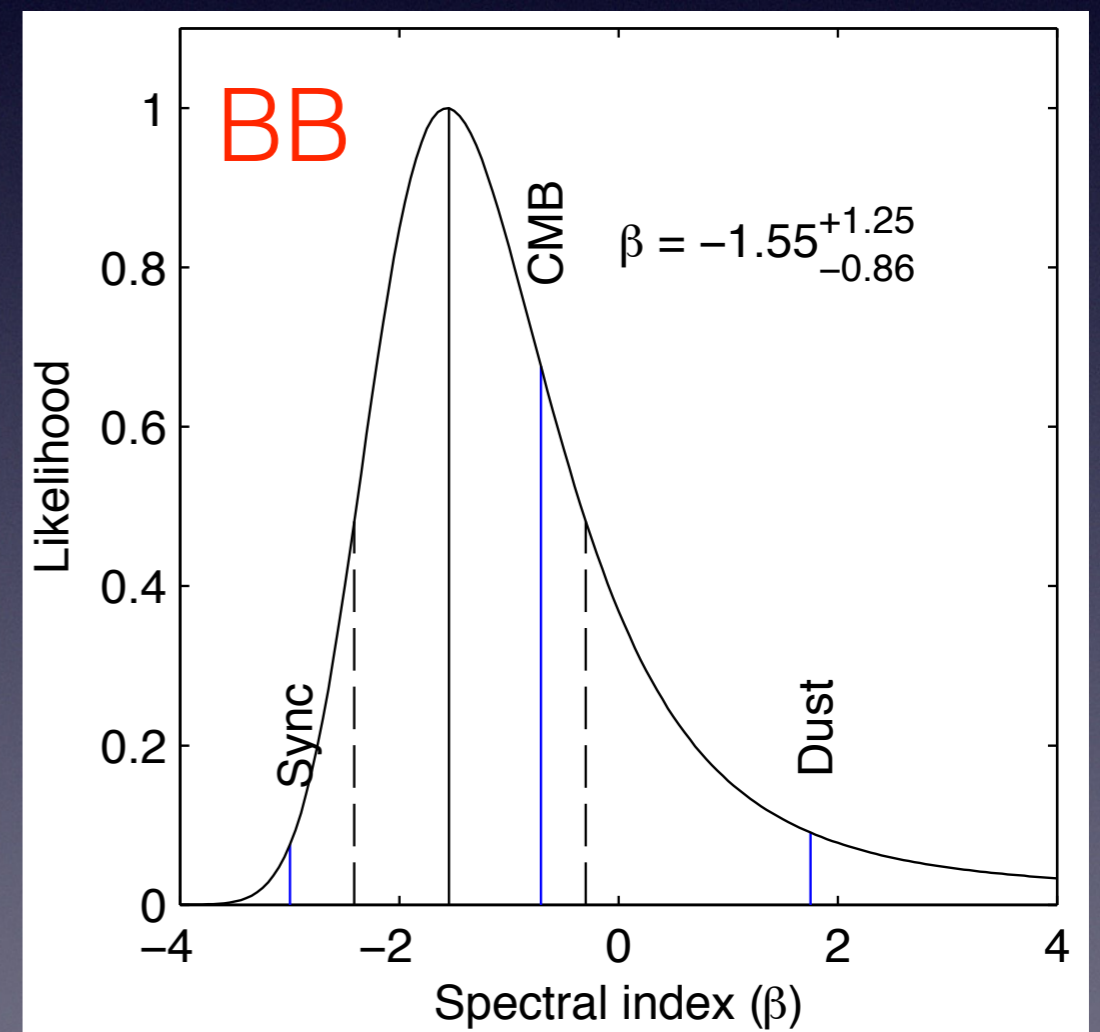


# Spectral Index Constraint



- Constrain BB signal color with  $B_{2150} \times B_{1100}$
- If **dust**, expect little correlation
- If **synchrotron**, expect bright correlation

- **Consistent with CMB**, disfavor benchmark dust and sync models at **2.2/2.3 $\sigma$**



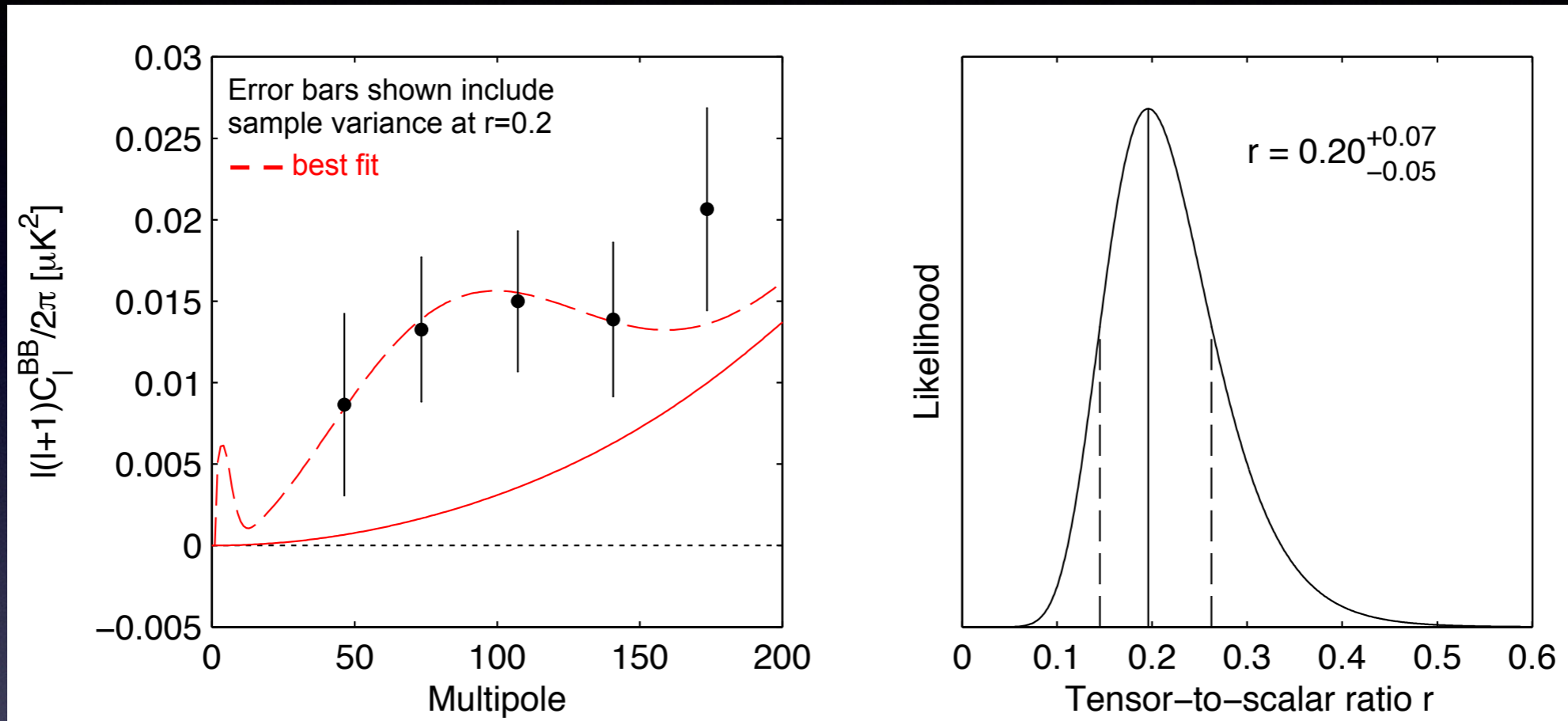
*Antenna temperature frequency power law*

# What could this be?

- Instrumental systematics?
- Galactic foregrounds?
- **Cosmology**

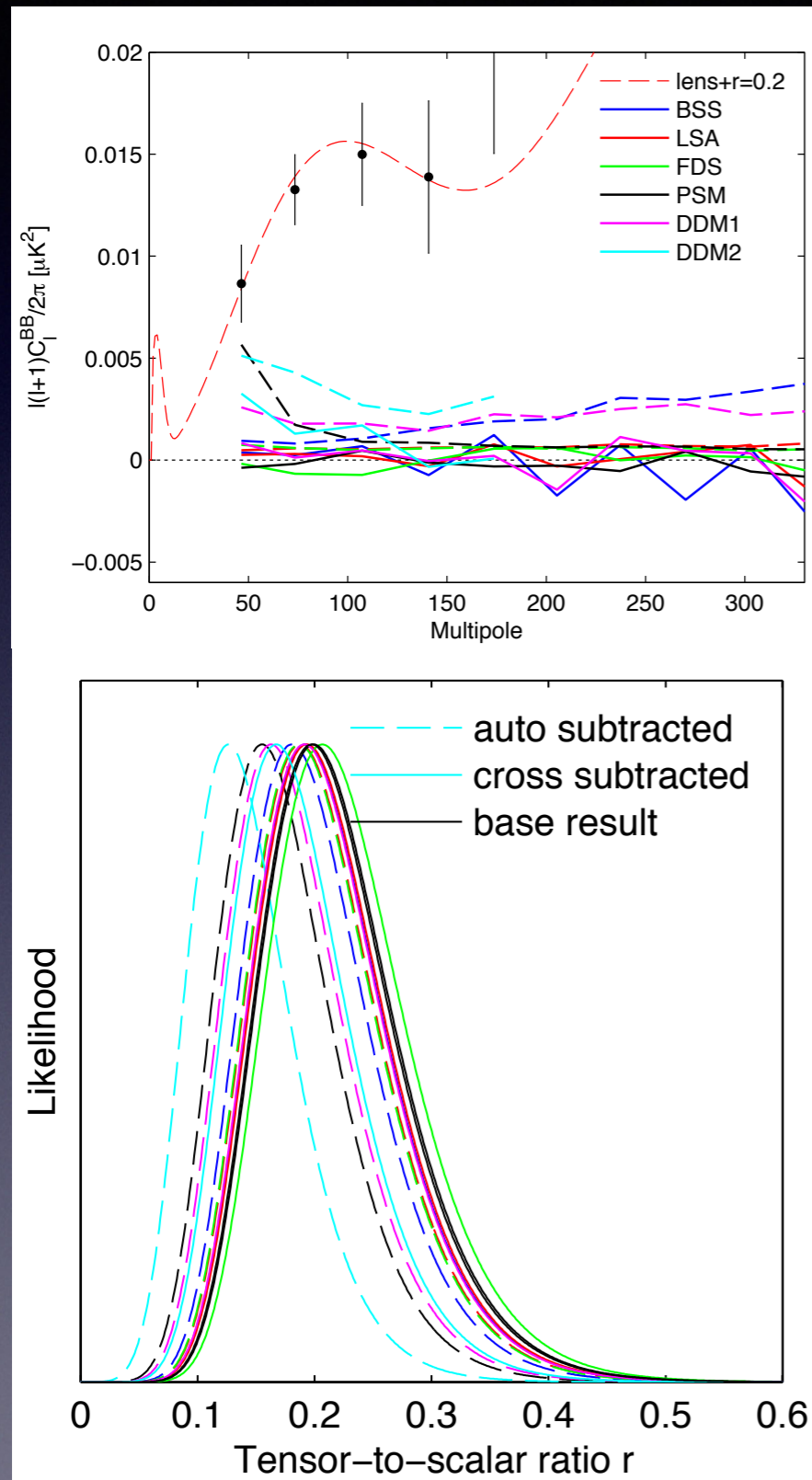


# Constraint on Tensor/Scalar Ratio



- Best-fit  **$r=0.20$**  (PTE of fit 0.9)  
*Consistent with large-field, GUT-scale inflation*
- $r=0$  disfavored at  **$7.0\sigma$**  (PTE  $3.3 \times 10^{-12}$ )
- **Sample variance dominated** -> *Need more sky!*

# Effect of Foregrounds



- Foregrounds could contribute small amount of observed BB
- Total power spectrum does not look like foreground expectations
- Subtracting DDM2 cross gives  $r_{sub} = 0.16^{+0.06}_{-0.05}$
- ... still disfavors  $r=0$  at  **$5.9\sigma$**
- Dust contributes most in the first band power. Deweighting this bin would give less deviation from our base result

# What next?

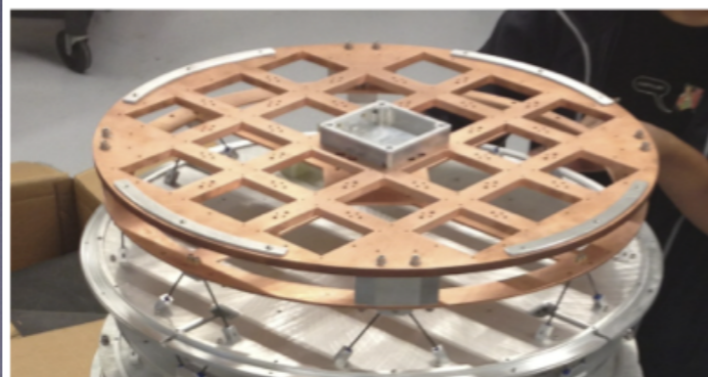
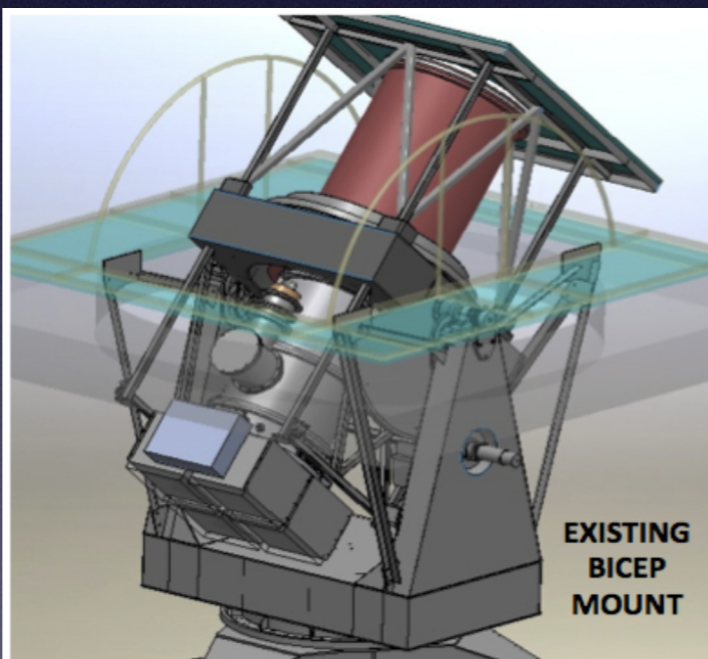
## KECK ARRAY

- South Pole, 2011 - 2016
- 2011: 1536 TESs @ 150 GHz
- 2012-13: 2560 @ 150 GHz
- 2014 upgrade:
  - 1536 @ 150 GHz
  - 576 @ 100 GHz



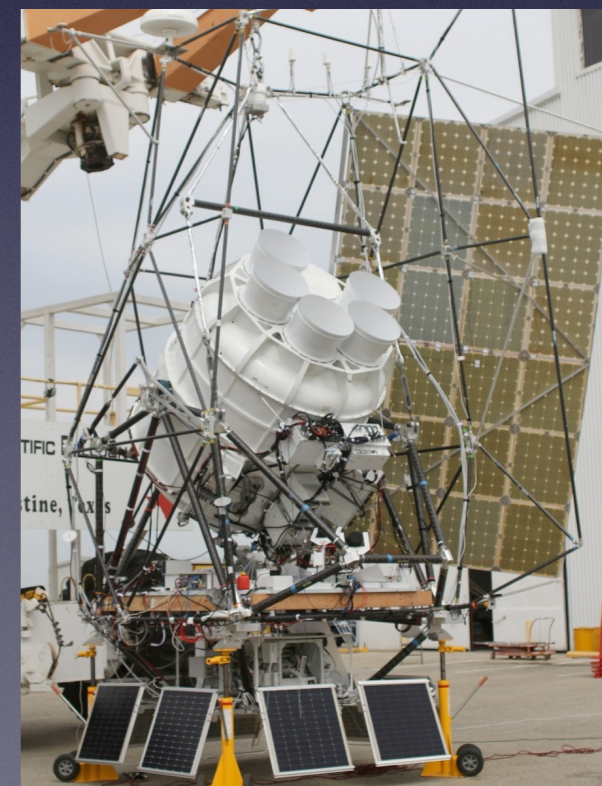
## BICEP3

- South Pole, 2015-16
- 2560 TESs @ 100 GHz

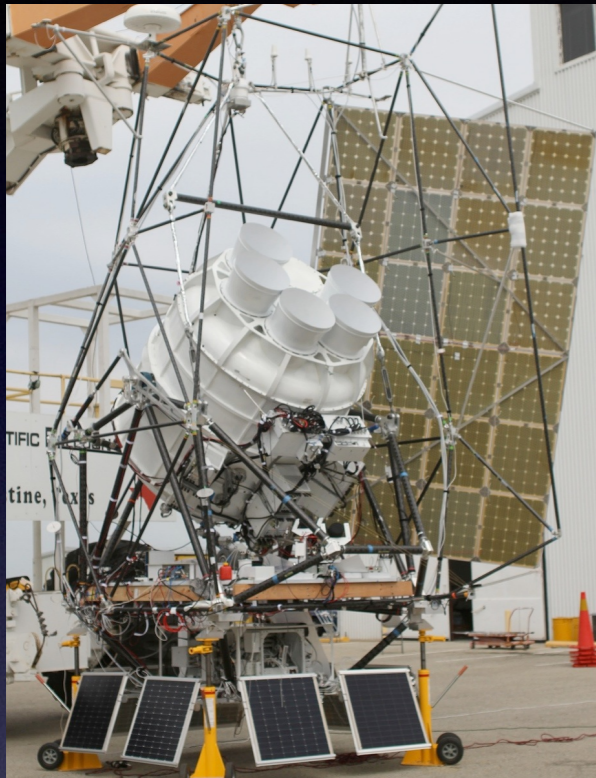


## SPIDER

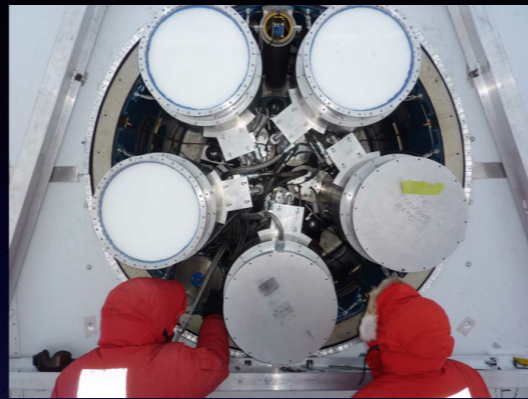
- Long-duration balloon 2014
- Large (~10%) sky coverage
- Half-wave plate
- 2400 TESs
  - 1536 @ 150 GHz
  - 864 @ 100 GHz



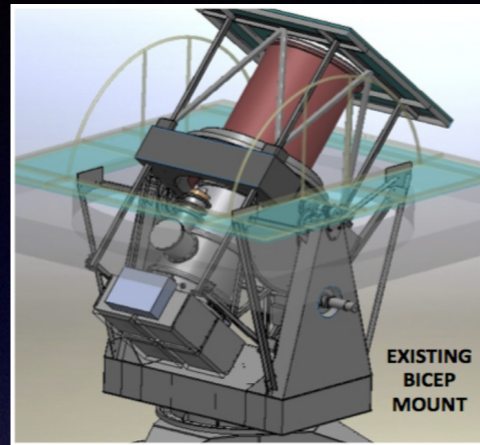
# A Rich Field



SPIDER



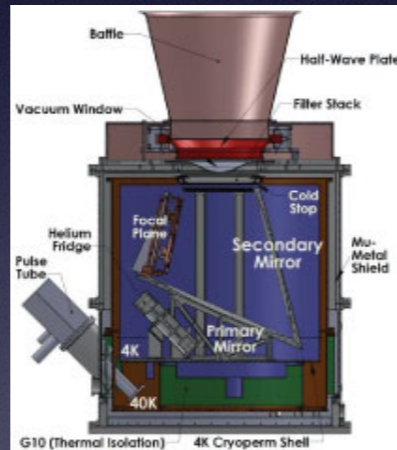
Keck Array



BICEP3



EBEX



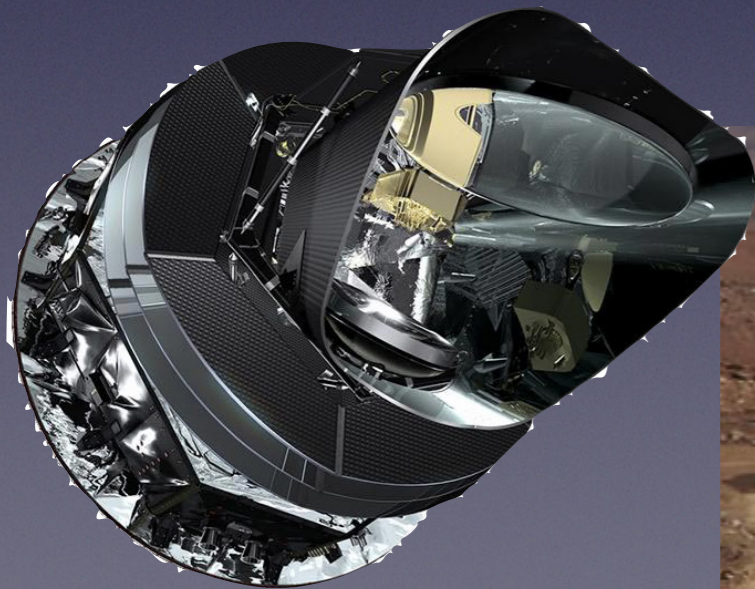
ABS



SPTpol



POLARBEAR

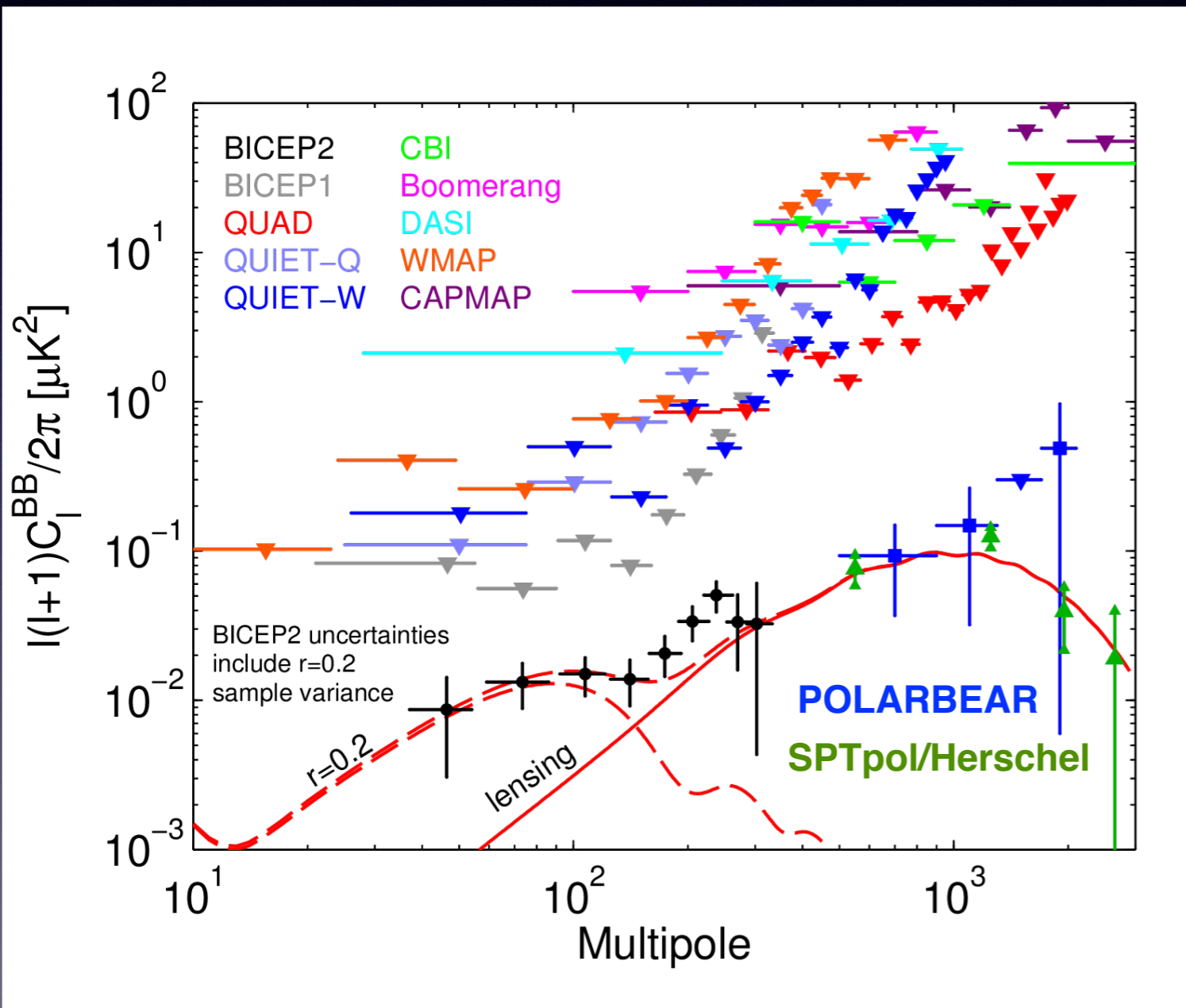


Planck



ACTpol

# Conclusions



- BICEP2 observes **5.3 $\sigma$  excess** above lensed- $\Lambda$ CDM;  $r=0$  disfavored at **7 $\sigma$**  (no foreground subtraction)
  - Extensive studies disfavor systematic error as origin
  - Foregrounds do not appear to constitute the bulk of the signal
- Consistent with expectations for primordial gravitational waves from GUT-scale inflation
- The era of B-mode cosmology has begun!